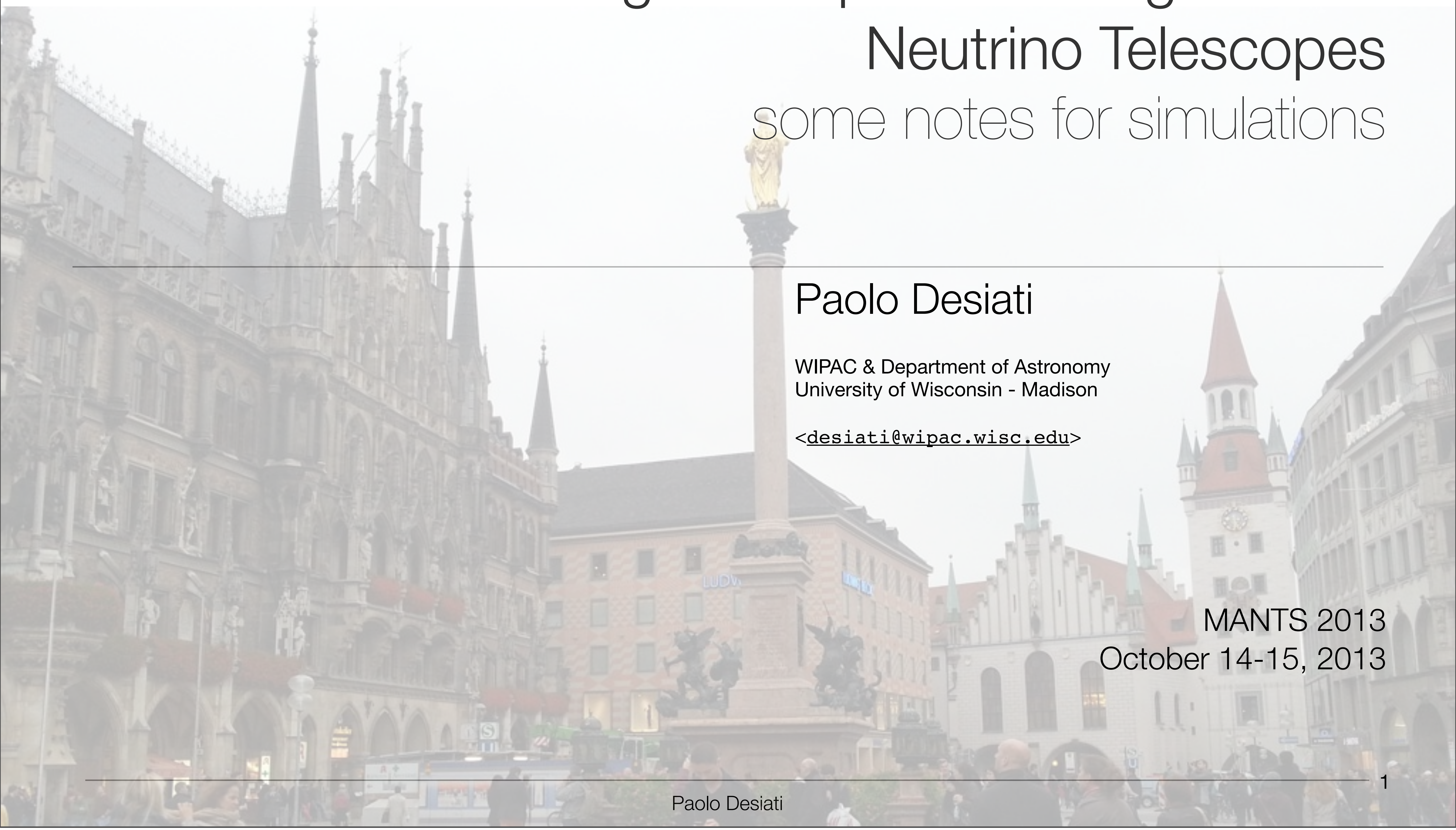
Wisconsin IceCube Particle Astrophysics Center

Bayerische Akademie der Wissenschaften

Understanding Atmospheric Background in Neutrino Telescopes

some notes for simulations



Paolo Desiati

WIPAC & Department of Astronomy
University of Wisconsin - Madison

<desiati@wipac.wisc.edu>

MANTS 2013
October 14-15, 2013

cosmic rays

produce background for neutrino searches but it's also a signal.

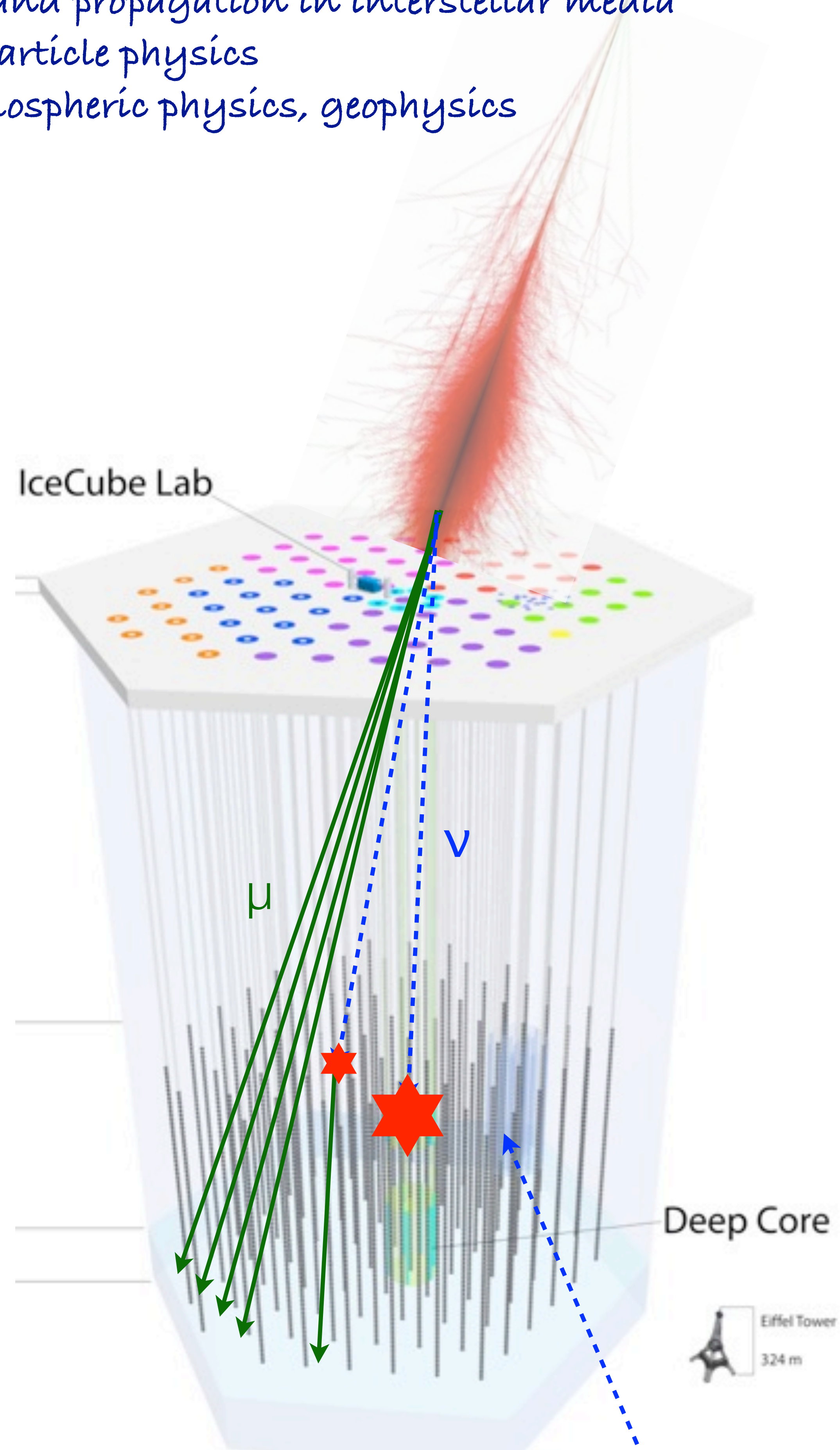
are linked to astrophysics and propagation in interstellar media
their interactions involve particle physics
their by-products probe atmospheric physics, geophysics

primary cosmic rays
spectrum
mass composition

atmospheric target
density & temperature profile

hadronic interaction models
primary interaction
secondary interactions
interaction cross sections
fragmentation & hadronization
heavy quarks
(non) perturbative processes

propagation
electromagnetic component
penetrating component



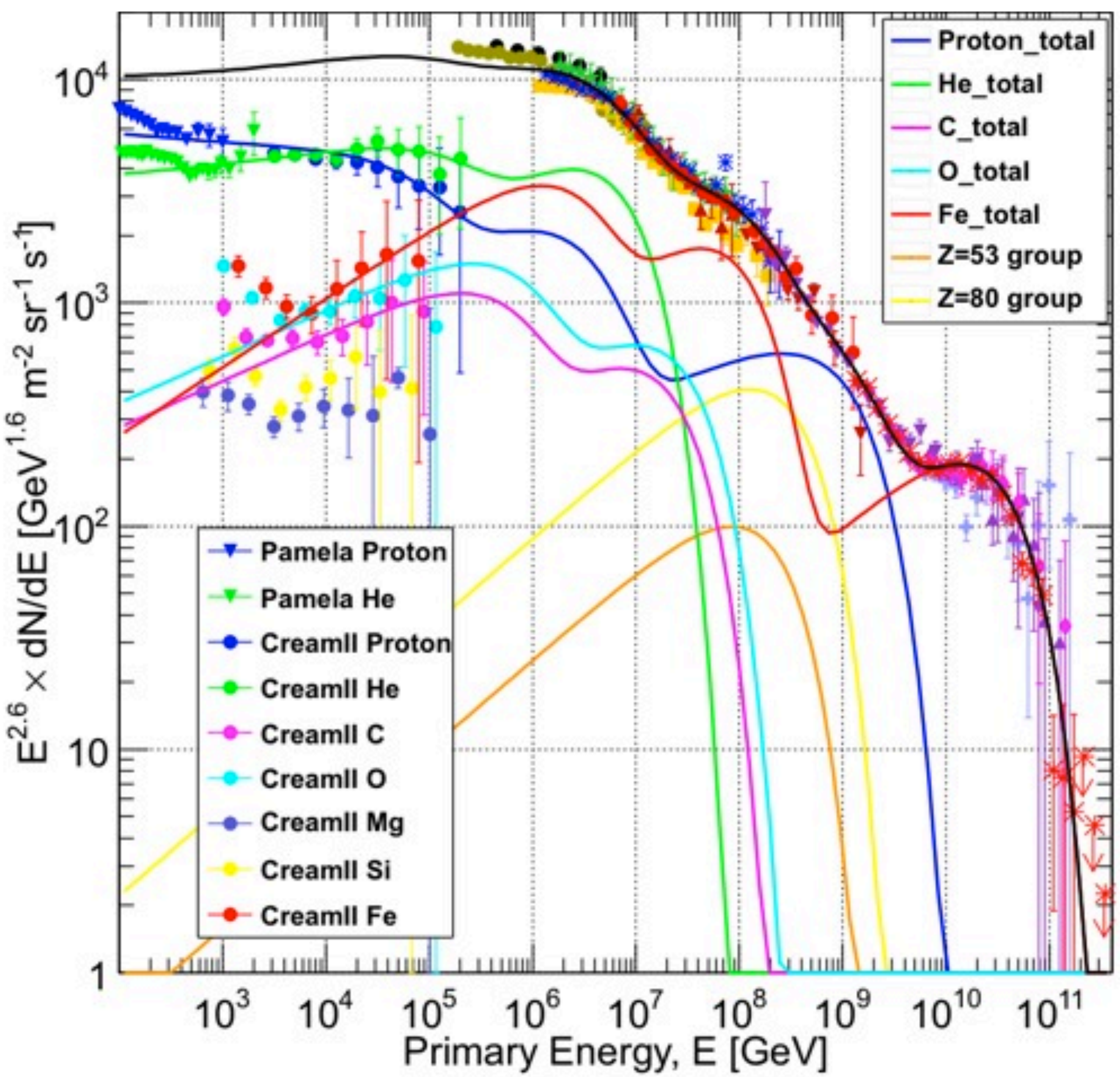
primary cosmic rays
spectrum & composition

- cosmic ray all particle spectrum observed in wide energy range
- composition directly observed < 100 TeV
- statistically estimated at high energy
- large uncertainties on mass composition
- affect **atmospheric lepton uncertainties**

direct measurements

indirect measurements

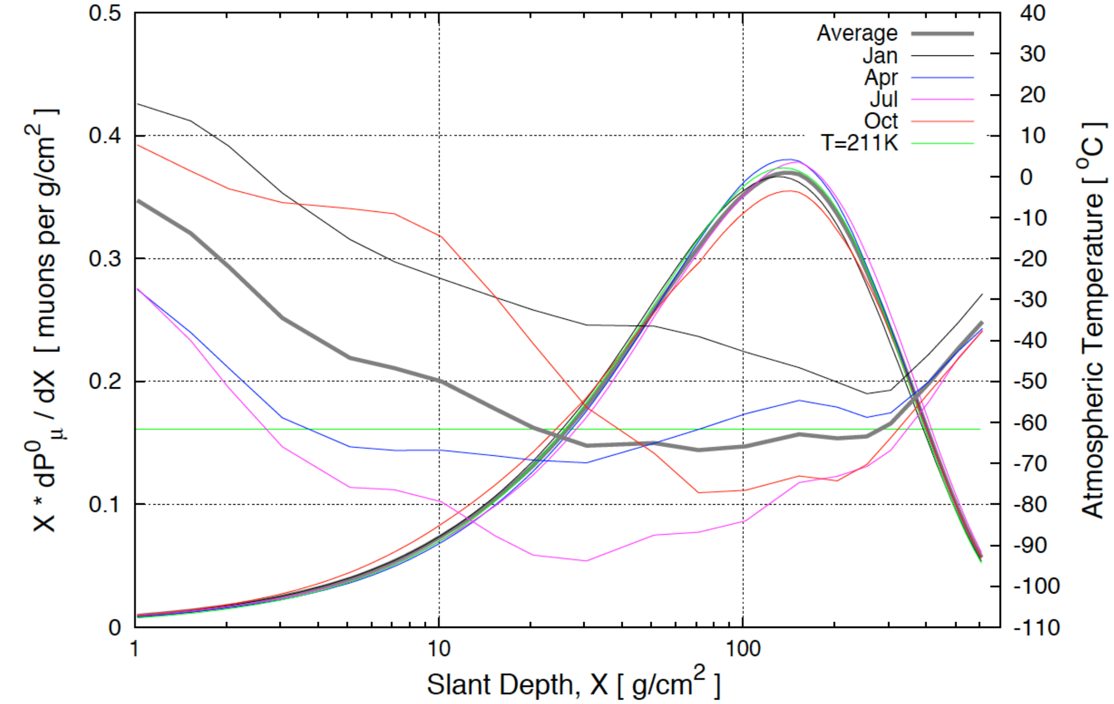
Gaisser, Stanev & Tilav, 2013 - arXiv:1303.3565

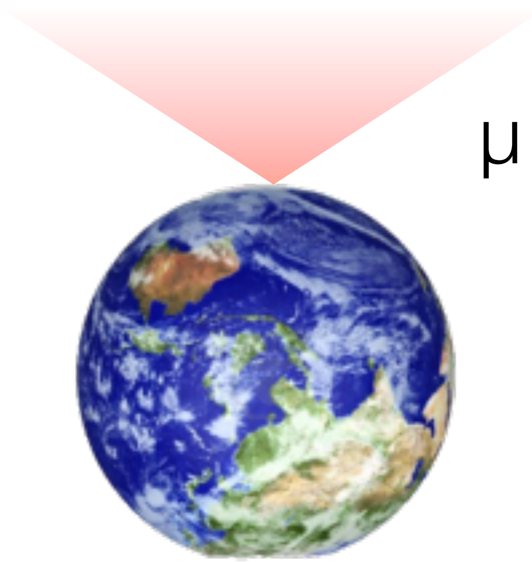


atmospheric target

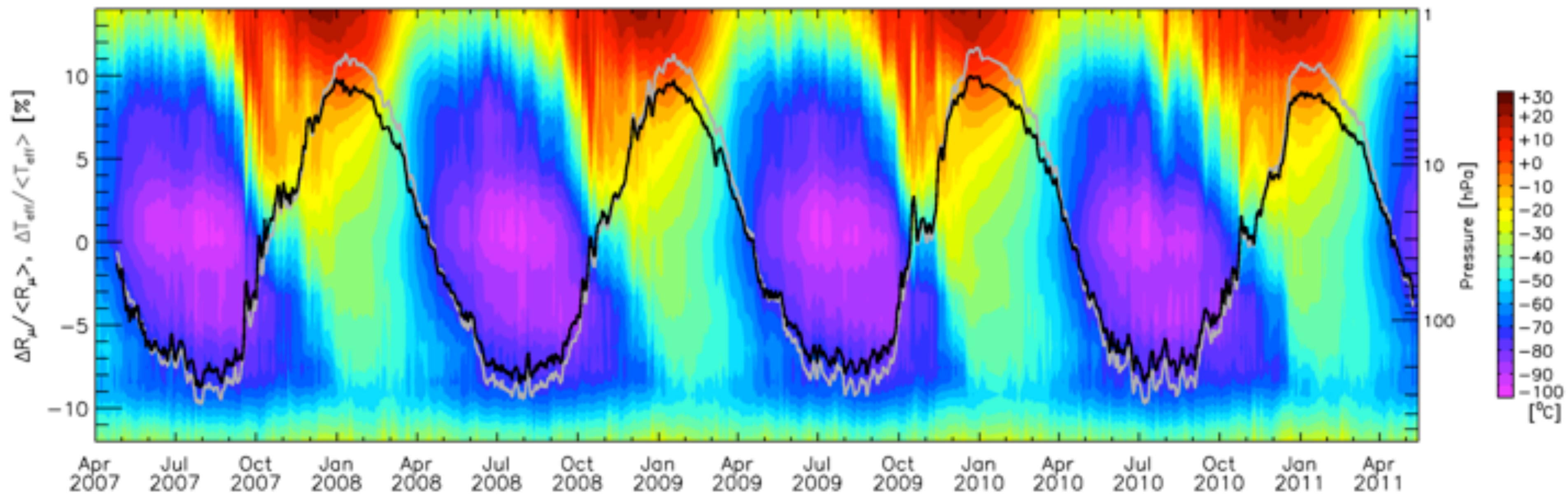
- cosmic ray target changes properties over the seasons

μ production spectrum & atmospheric temperature



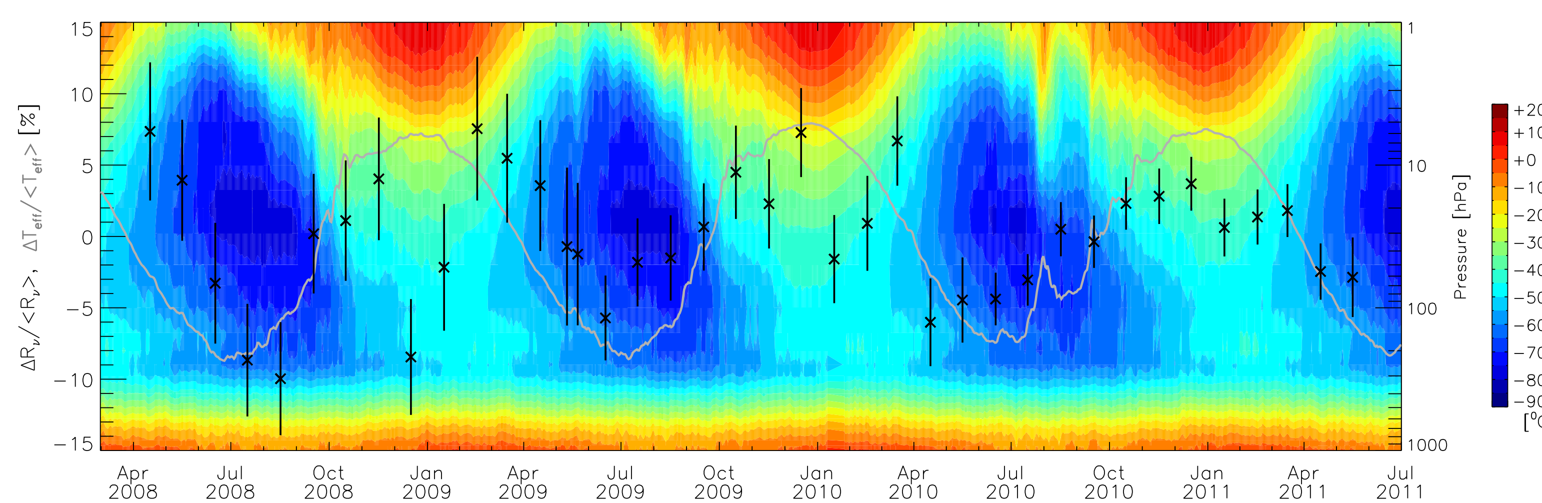


ICRC 2011



μ multiplicity - ICRC 2013

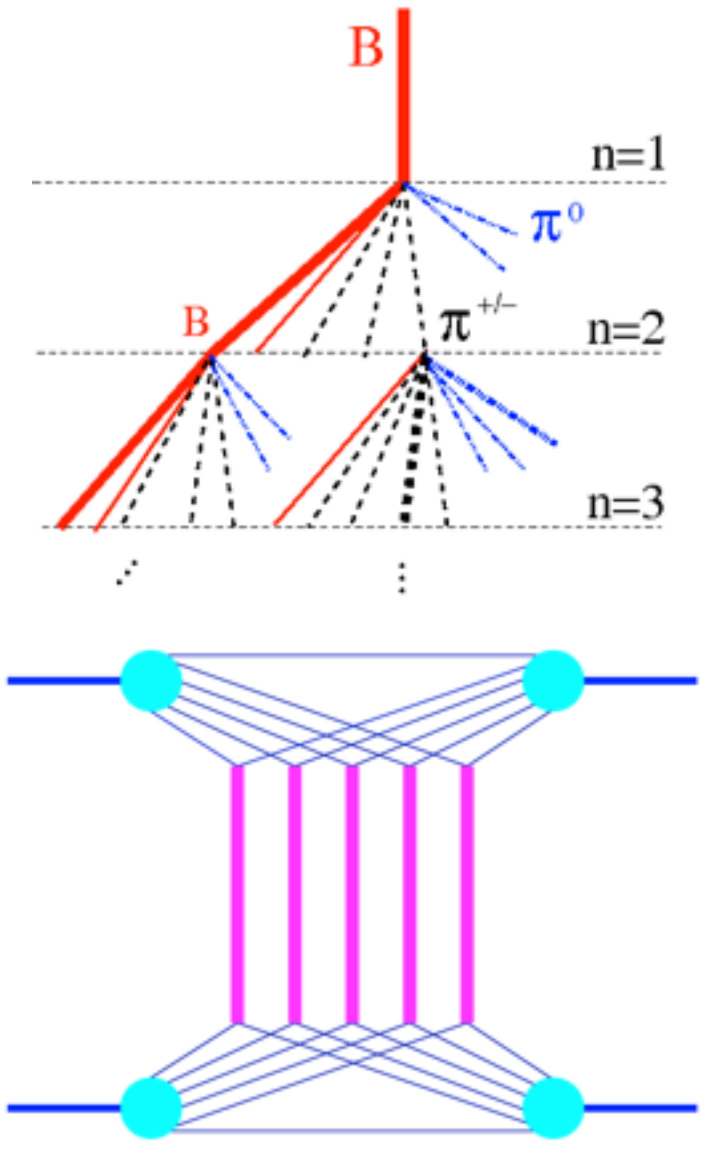
ICRC 2013



source of uncertainties

hadronic interactions

source of uncertainties in CR physics



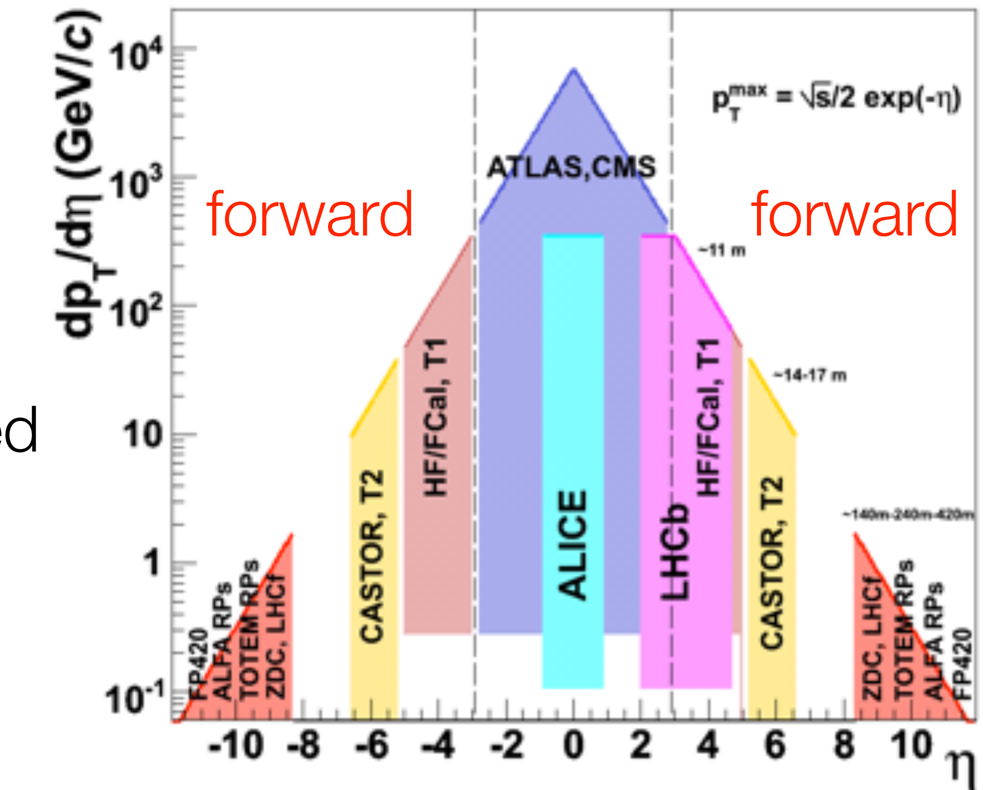
- CR showers dominated by **soft component with small p_T** (*non-perturbative QCD*)
- **hard component with high p_T** with heavy quarks (pQCD)
- **phenomenological** descriptions of hadronic interactions with minijet production for hard component
- **models** to describe soft/hard **interactions** in **forward region** & **extrapolated to high energy**

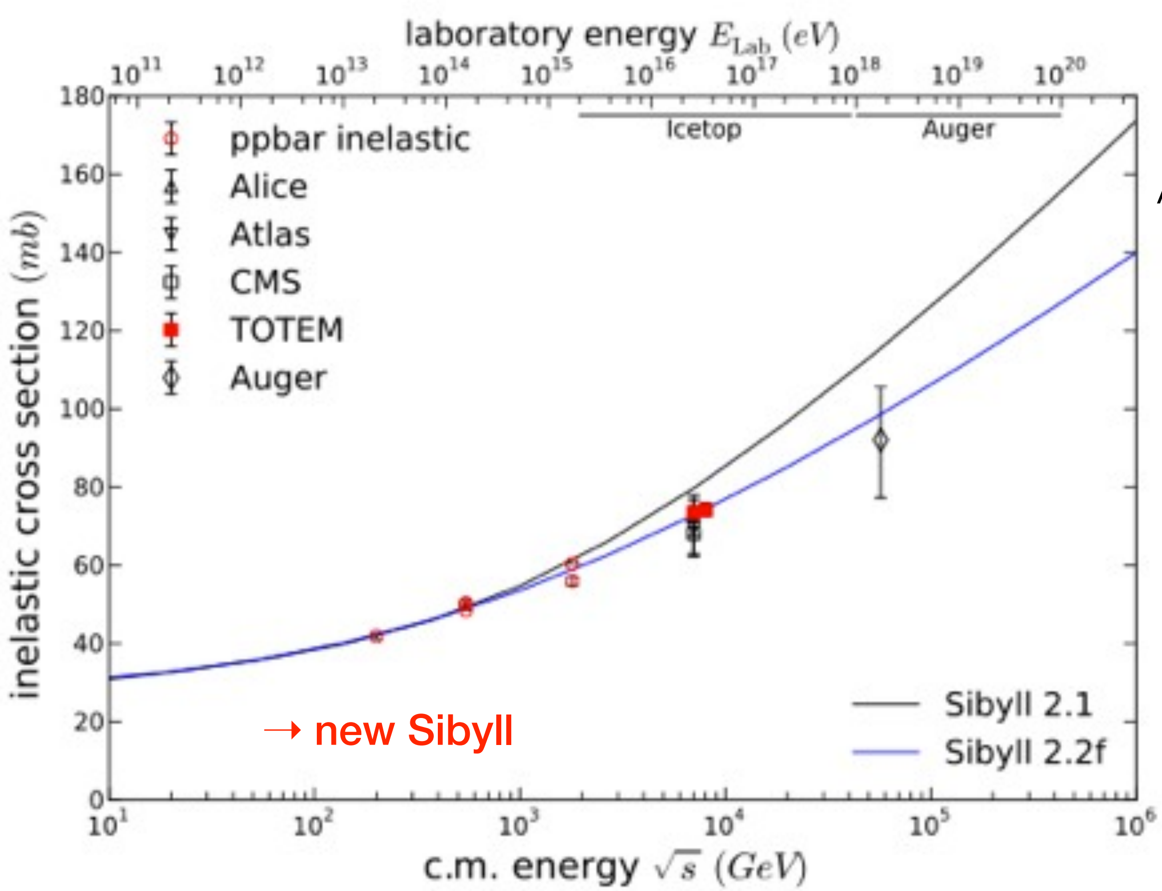
- **interaction models** from accelerators, **extrapolated** to forward region at high energy

hadronic interactions

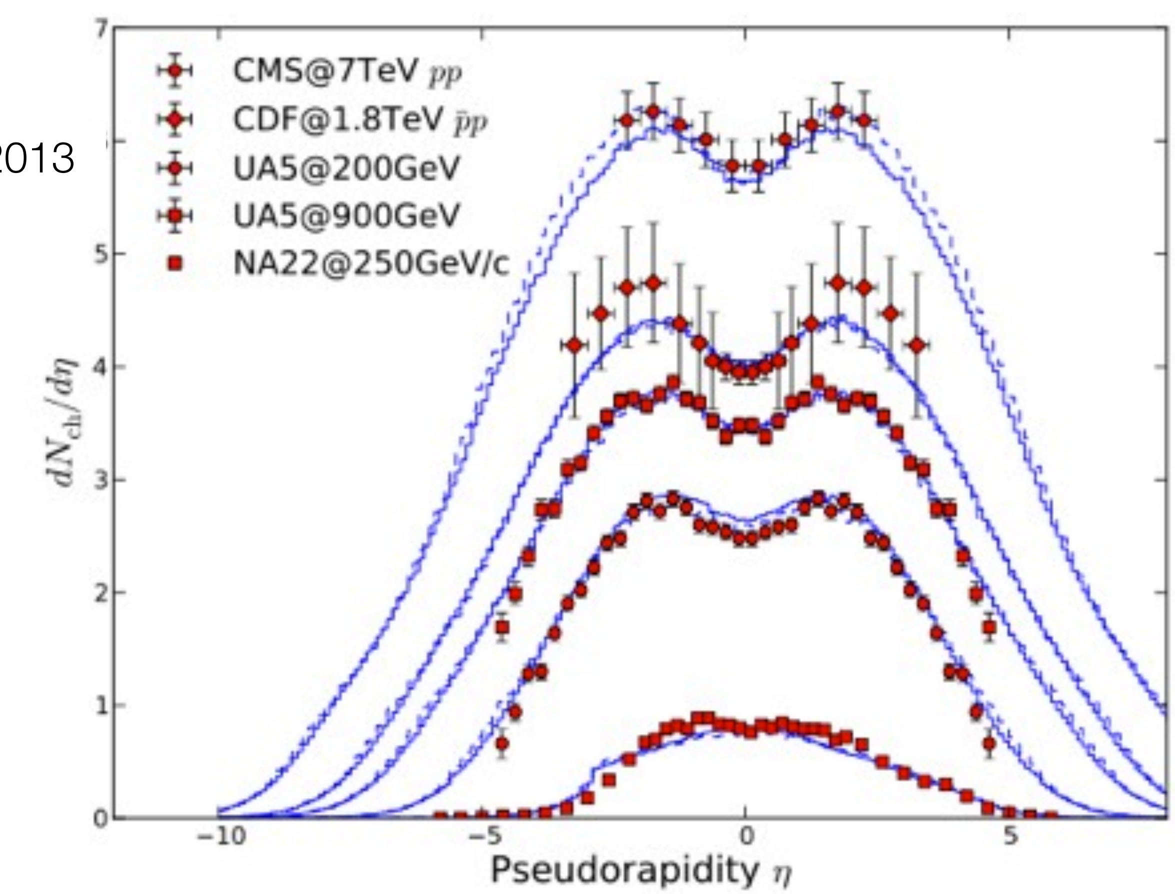
- **forward region** the most relevant in cosmic rays
- models **tuned** to accelerator measurements and extrapolated

transverse momentum p_T
vs
pseudo-rapidity



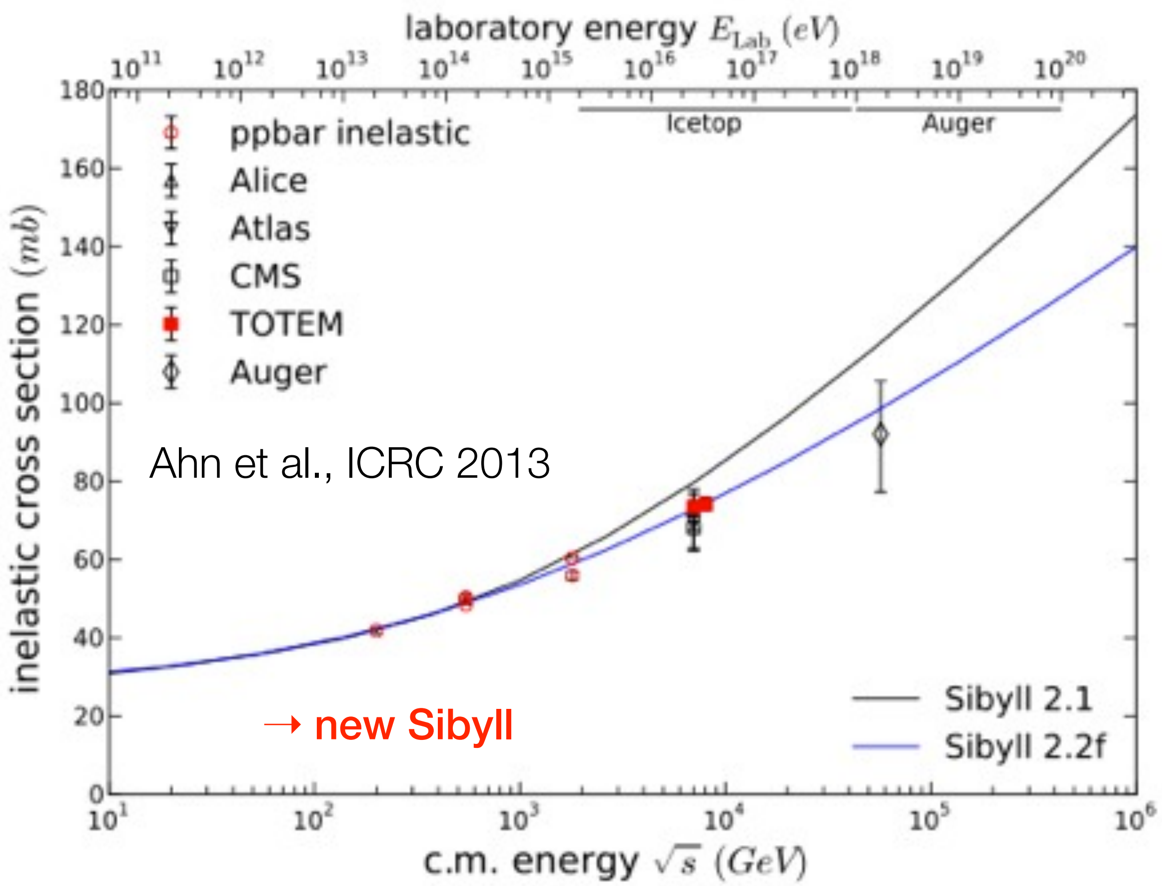


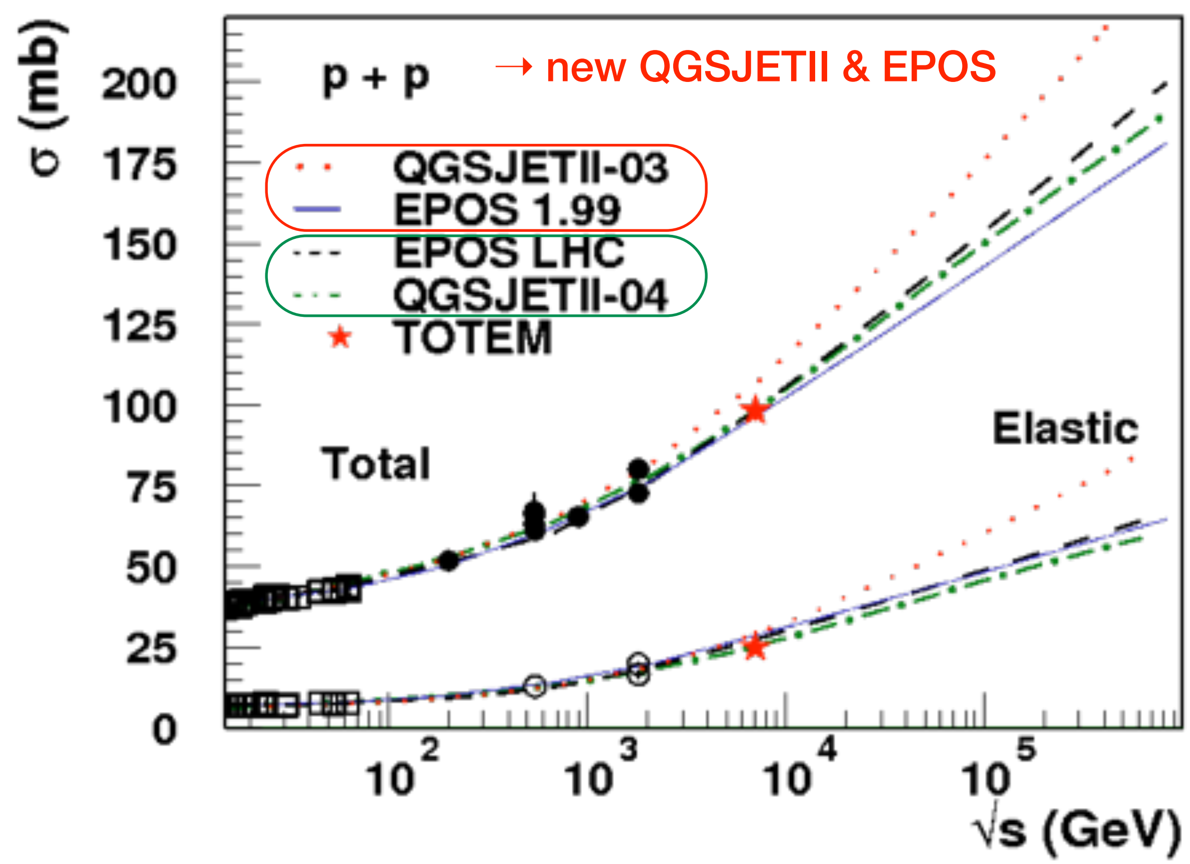
Ahn et al., ICRC 2013



hadronic interactions

- **forward region** the most relevant in cosmic rays
- models **tuned** to accelerator measurements and extrapolated
- LHC experiments (e.g. TOTEM, LHCf) starting to fill the relevant parameter space

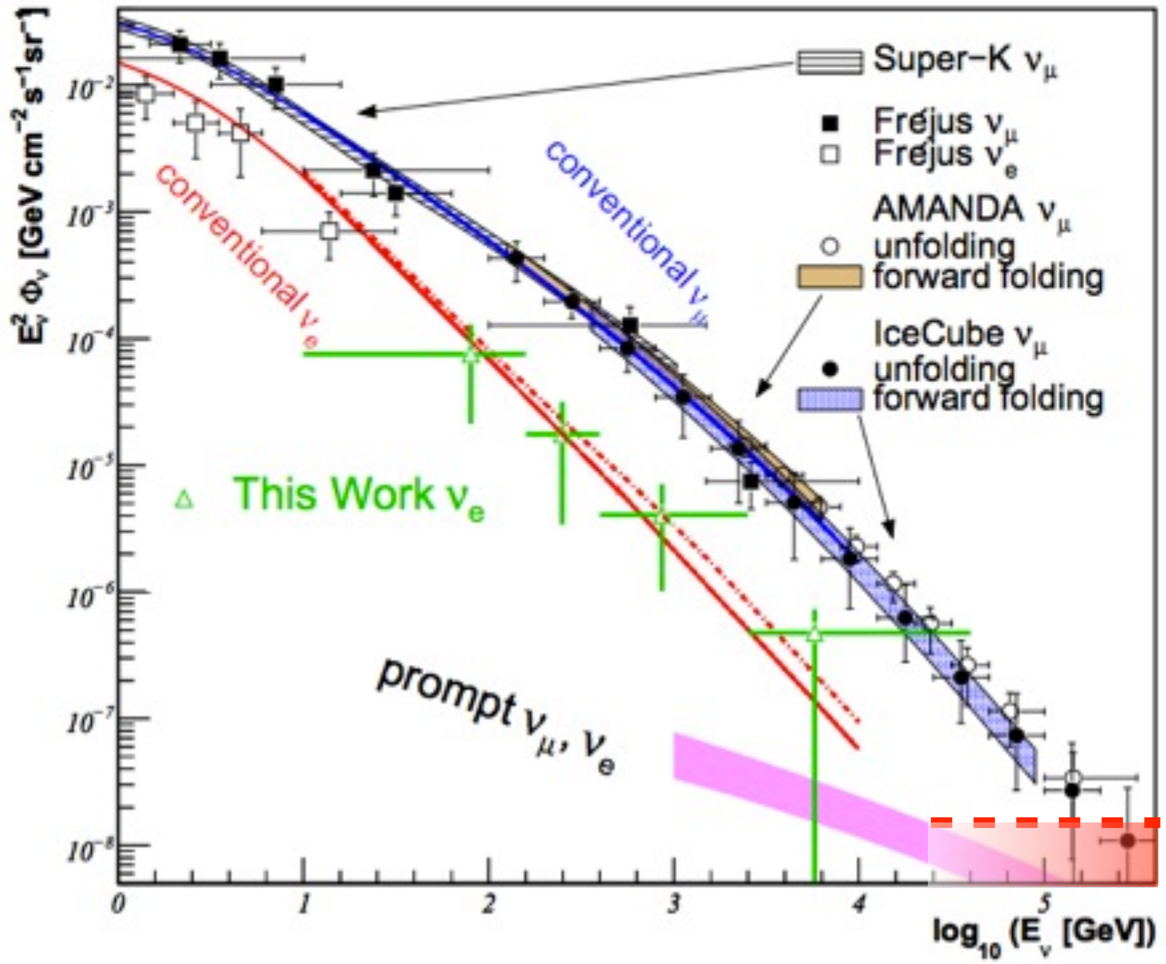




IceCube Coll., PRL 110, 151105 (2013)

atmospheric neutrinos
high energy and heavy quarks

- **neutrino telescopes** searching for high energy astrophysical neutrinos (*point to origin of CR*)
- **atmospheric neutrinos** as **irreducible background** at high energy where **heavy quark processes** are involved
- production of **hyperons** and particles with **charm** affected by increasing uncertainties
- **CORSIKA** numerical simulations



$$\phi_\nu(E_\nu) = \phi_N(E_\nu) \times \left\{ \frac{A_{\pi\nu}}{1 + B_{\pi\nu}\cos\theta\, E_\nu/\epsilon_\pi} + \frac{A_{K\nu}}{1 + B_{K\nu}\cos\theta\, E_\nu/\epsilon_K} + \frac{A_{\mathrm{charm}\,\nu}}{1 + B_{\mathrm{charm}\,\nu}\cos\theta\, E_\nu/\epsilon_{\mathrm{charm}}} \right\}$$

$$A_{i\nu} = \frac{Z_{Ni} \times BR_{i\nu} \times Z_{i\nu}}{1 - Z_{NN}}$$

$$Z_{N\pi^\pm}(E) = \int_E^\infty \mathrm{d}E' \frac{\phi_N(E')}{\phi_N(E)} \frac{\lambda_N(E)}{\lambda_N(E')} \frac{\mathrm{d}n_{\pi^\pm}(E',E)}{\mathrm{d}E}$$

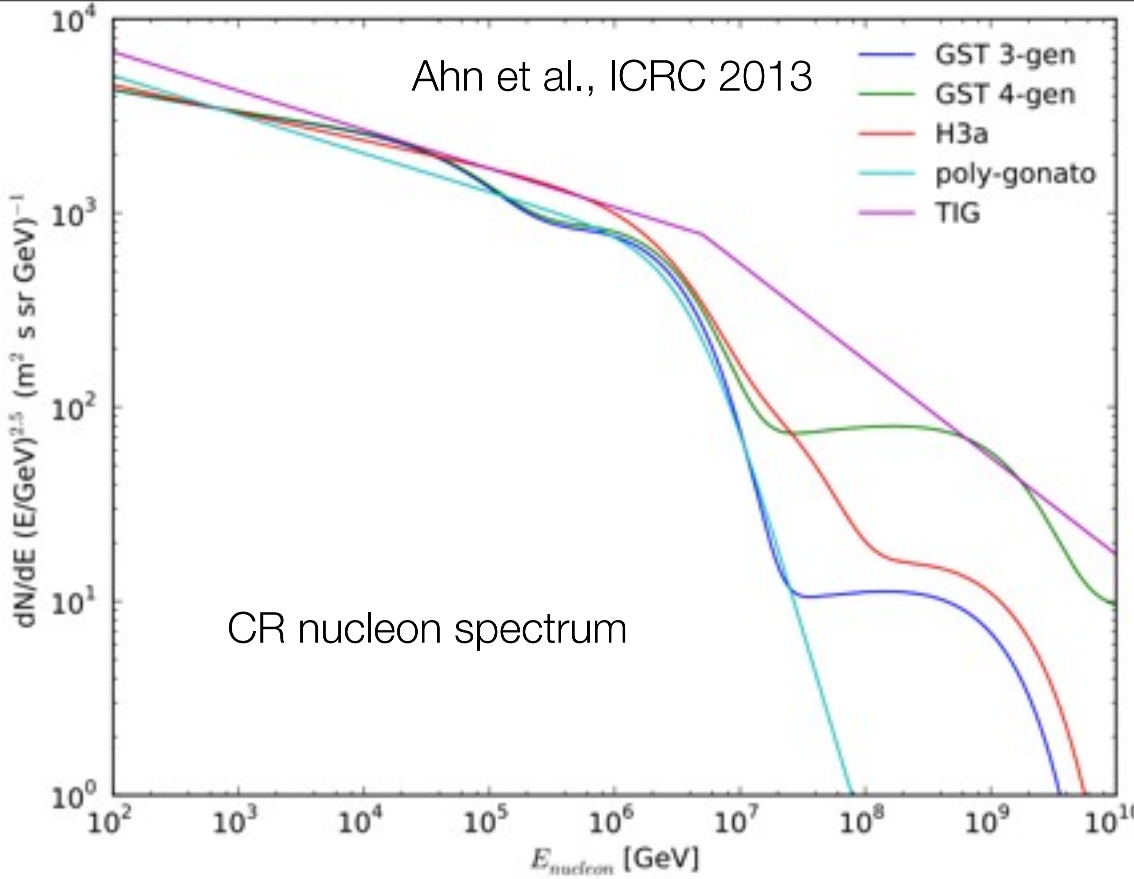
meson's characteristic energy

Particle (α):	$\pi^\pm$	$K^\pm$	K_L^0	Charm
ϵ_α (GeV):	115	850	205	$\sim 3 \times 10^7$

Paolo Desiati

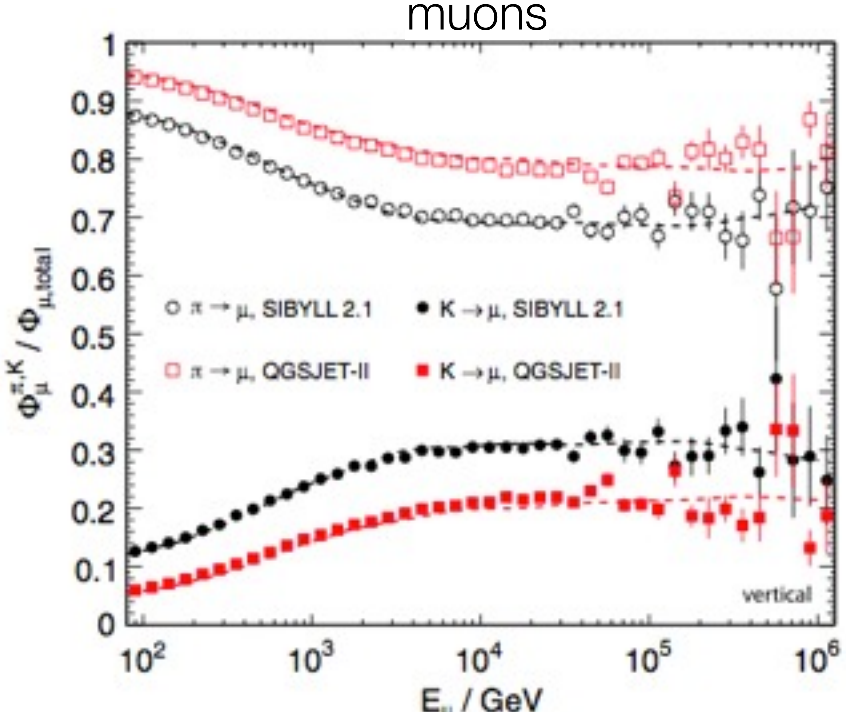
Monday, October 14, 2013

atmospheric neutrinos
high energy and heavy quarks

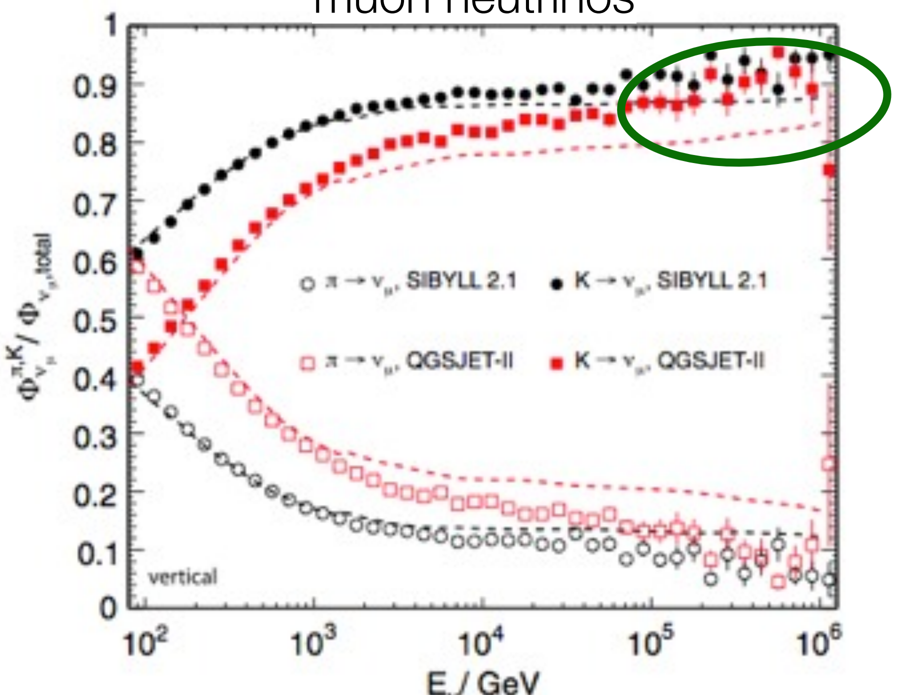


- large uncertainties in **cosmic ray composition** (nucleon spectrum) at high energy
- $K^{\pm}$ not same isospin group & K evolution equations **coupled**
- associated production p + Air → Λ + K^+ $\boxed{\mu^+/\mu^-,\ \nu_\mu/\bar{\nu}_\mu \quad \nu_e/\bar{\nu}_e}$

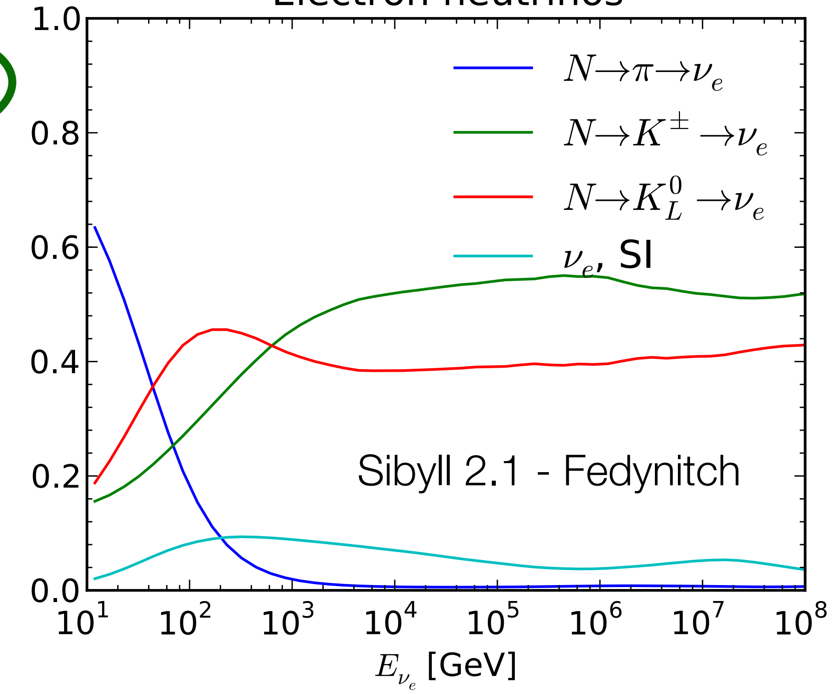
muons



muon neutrinos



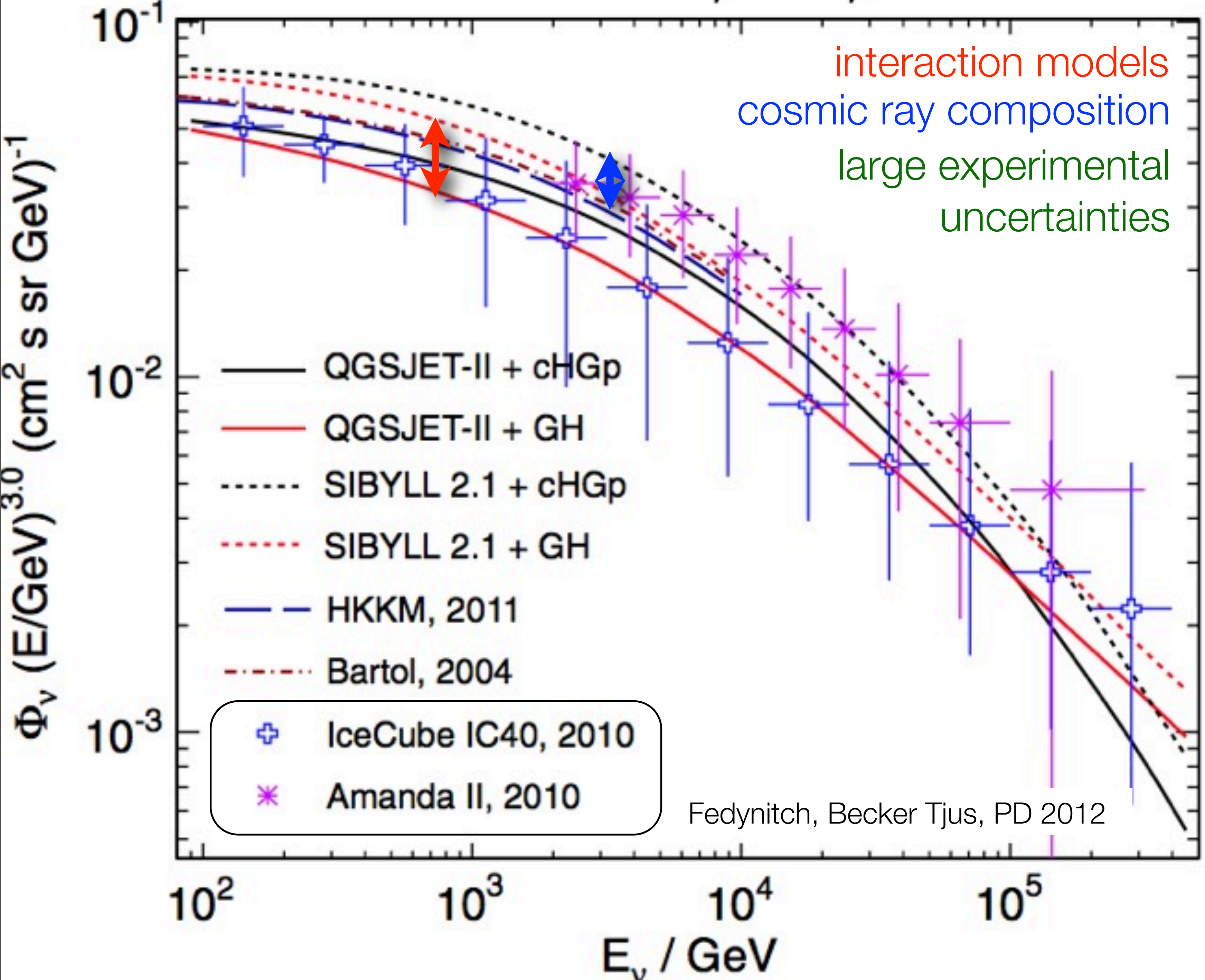
Electron neutrinos

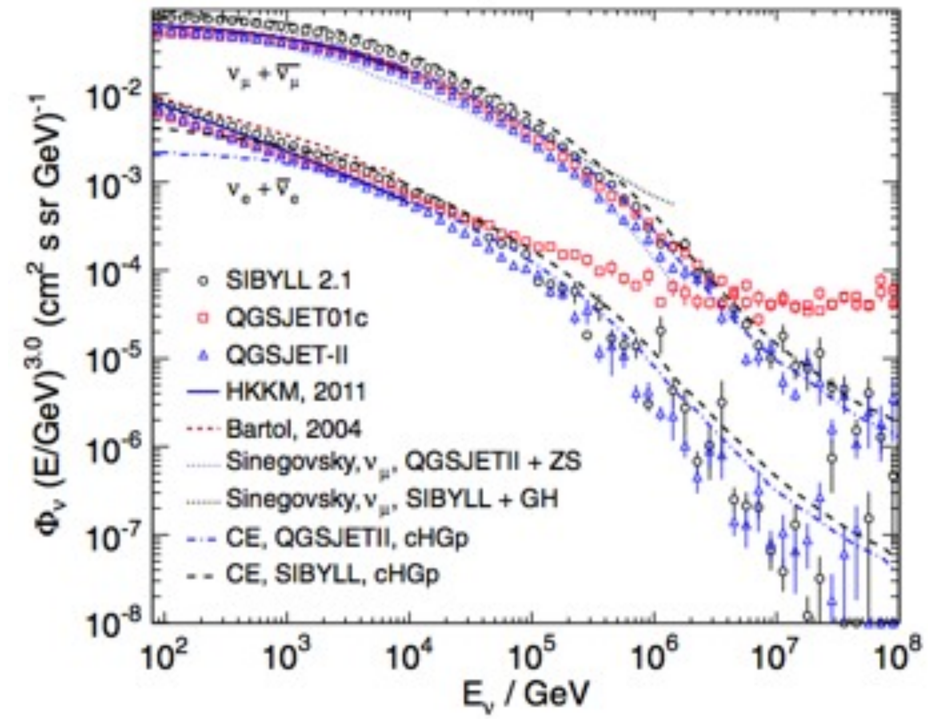


Fedynitch, Becker Tjus, PD 2012

atmospheric neutrinos
experimental observations

observed through-going $\nu_\mu + \bar{\nu}_\mu$





- extend **CORSIKA** production to neutrinos
- generation of muons & neutrinos with **consistent** primary **composition** and hadronic **interaction model**
- correlated **systematic effects** on backgrounds for physics analyses (zenith, energy, π/K, charm)

atmospheric neutrinos
charm production

- due to large quark mass, **perturbative QCD** can be used (hard component). However
  - significant charm production observed at $\sqrt{s}$ = 20 GeV
  - asymmetry in charm / anti-charm baryons (Selex Coll. 2002) → **intrinsic production**
- $|p\rangle = \alpha|uud\rangle + \beta|uudc\bar{c}\rangle + \ldots$: the **c-pair** produced in projectile fragmentation can recombine with valence quarks and with sea-quarks to **produce charmed hadrons**.

$p \rightarrow \Lambda_c^+ + \bar{D}^0$ ~ order $(m_s/m_c)^2$ (~1%) compared to $p \rightarrow \Lambda K^+$

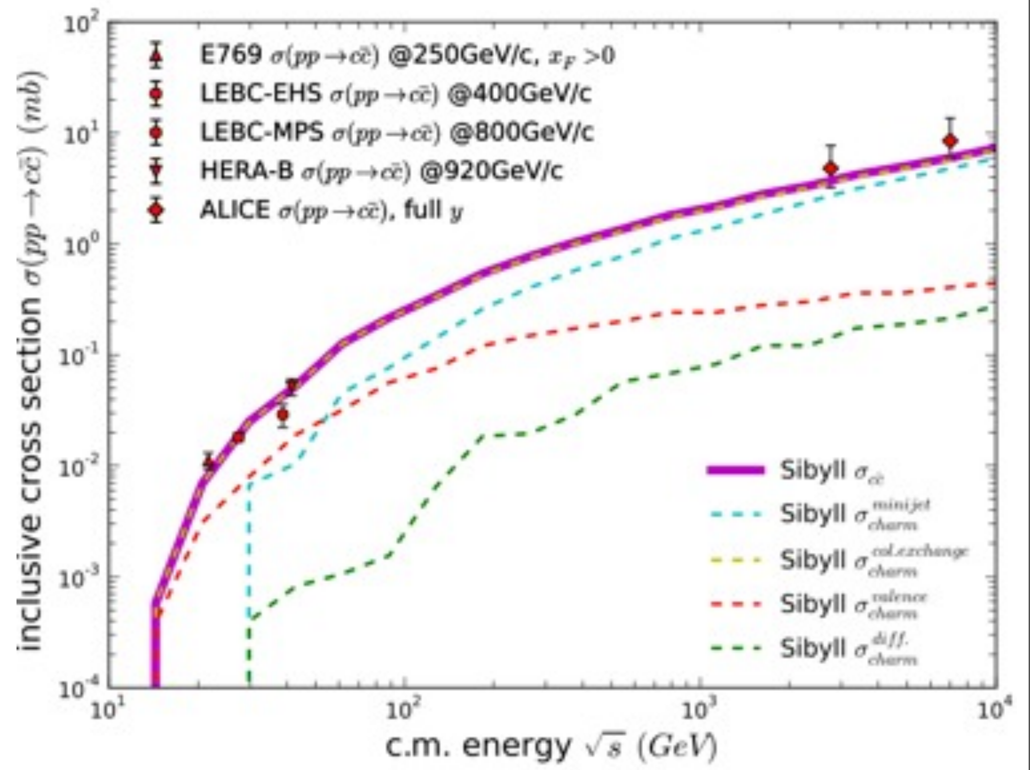
- inclusive D-meson spectrum dominated by intrinsic charm at high pseudo-rapidity & p_T

Lykasov+ 2012

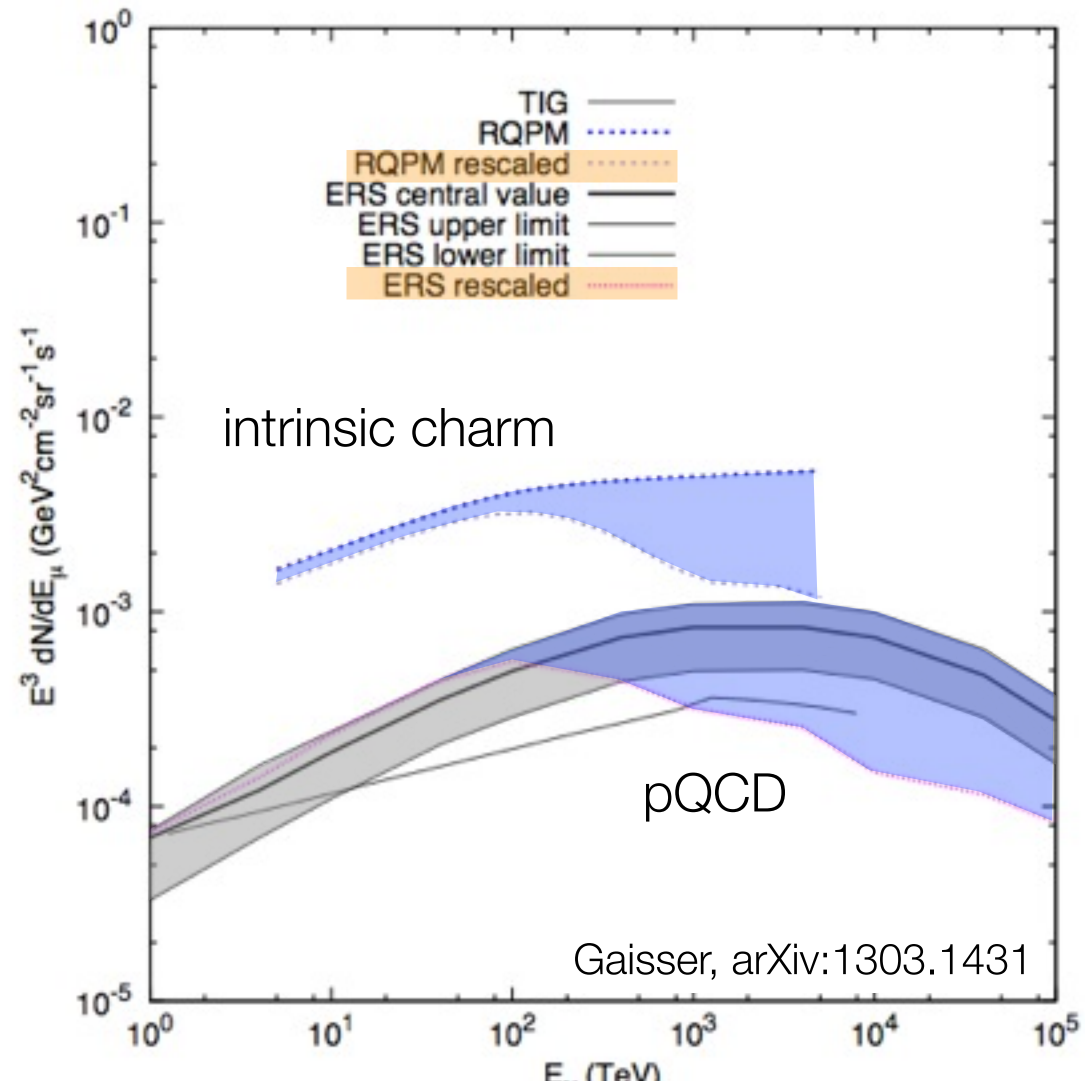
- steep cosmic ray spectrum might **enhance the effect of intrinsic production** of charm

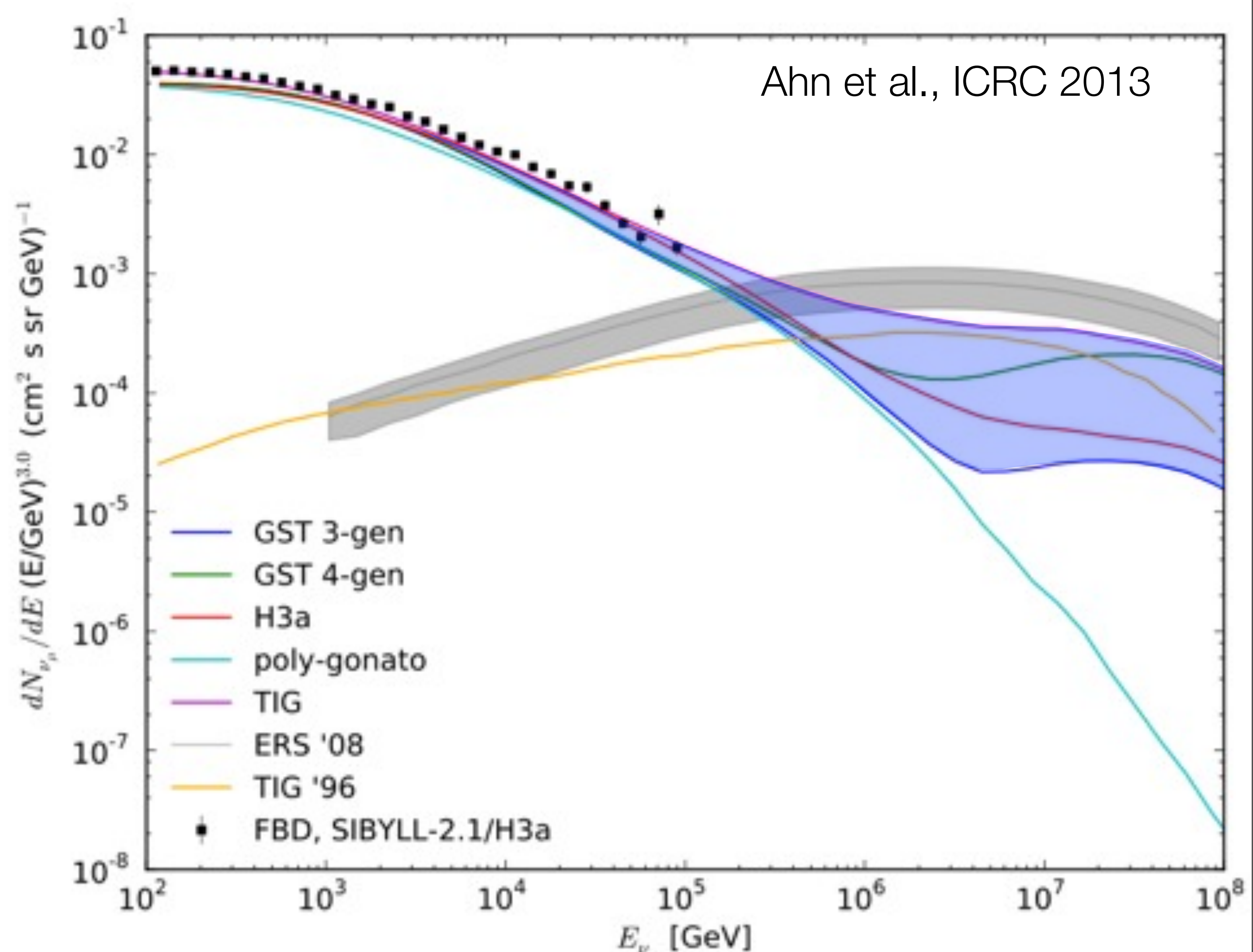
atmospheric neutrinos
charm production

- effect of charm production **models**
- effect of primary **cosmic ray spectrum**



Sibyll 2.2f - PRELIMINARY





atmospheric neutrinos

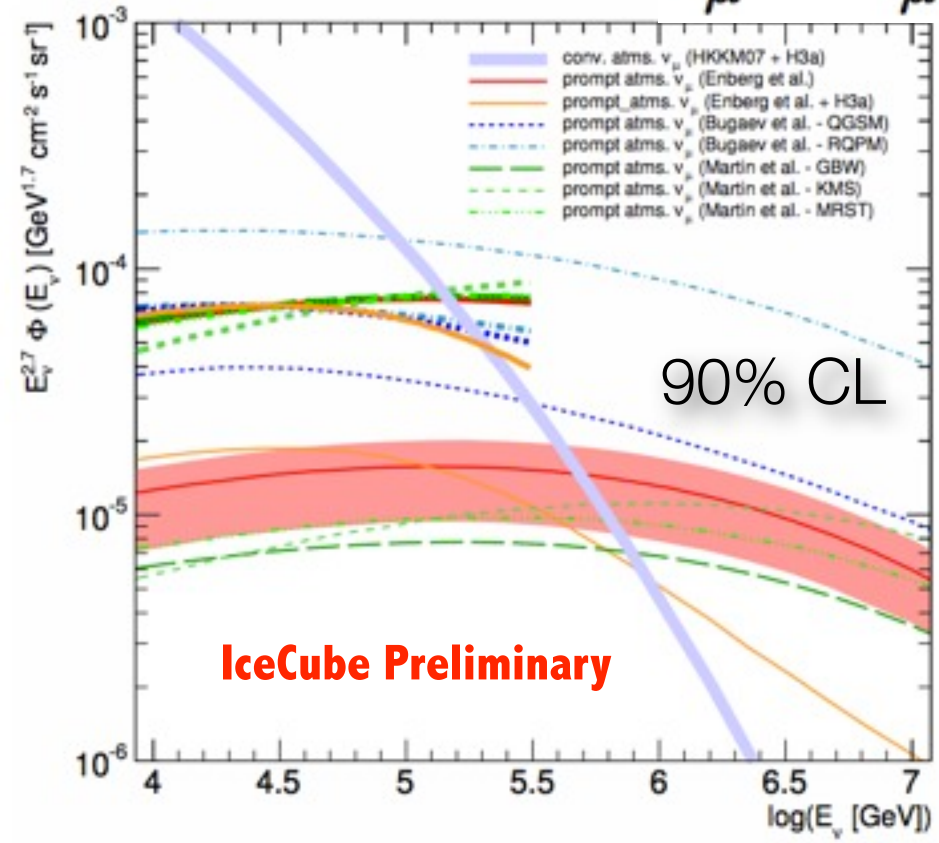
charm and astrophysical neutrinos

→ how easy is to measure an astrophysical signal ?
it depends on its spectrum

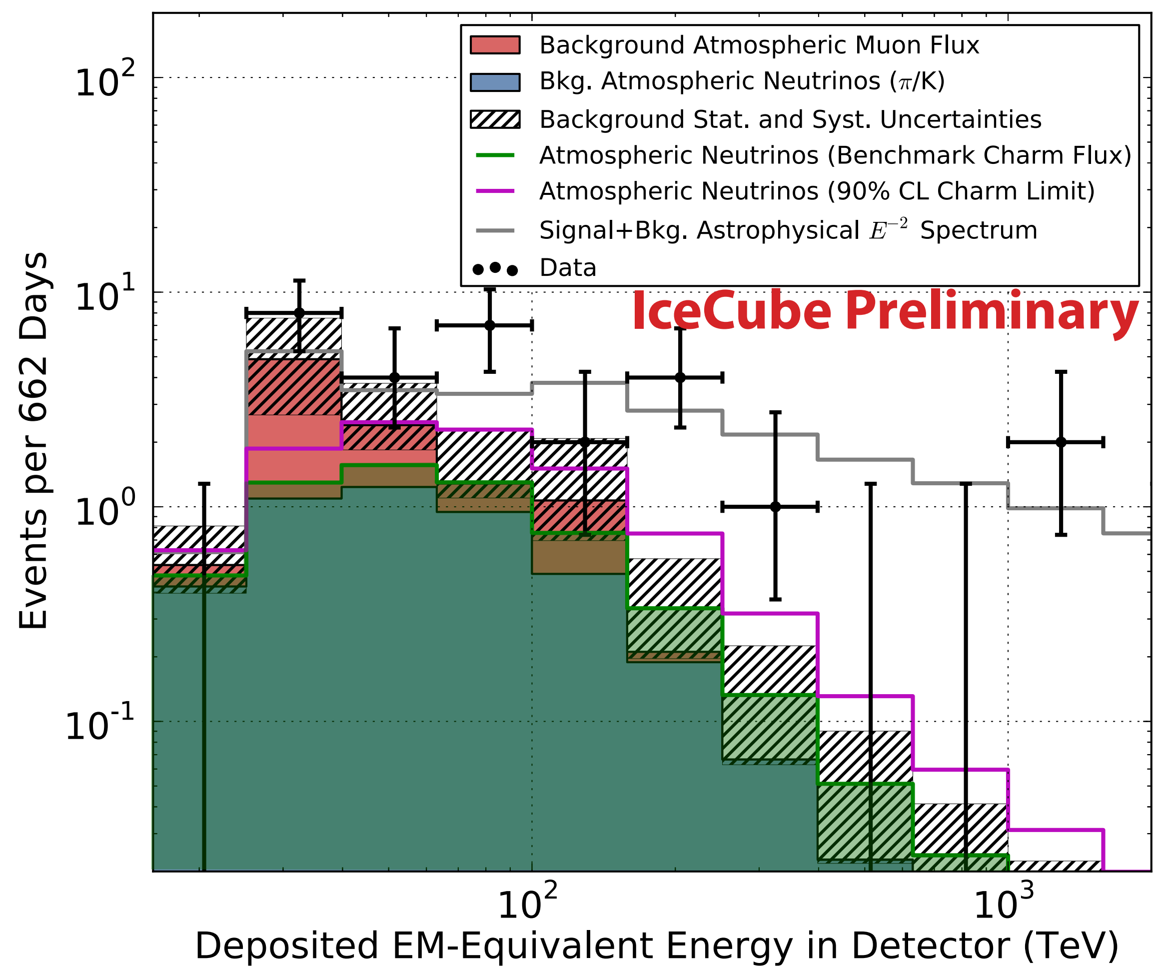
→ can neutrino telescopes ***observe*** neutrinos from charm ?
and constrain models ?
and break degeneracy between charm & astrophysics ?

→ assess model predictions in a comprehensive way

observed through-going up-ward $\nu_\mu + \bar{\nu}_\mu$



observed starting all-direction **all-flavor**



produce μ and ν simultaneously with CORSIKA

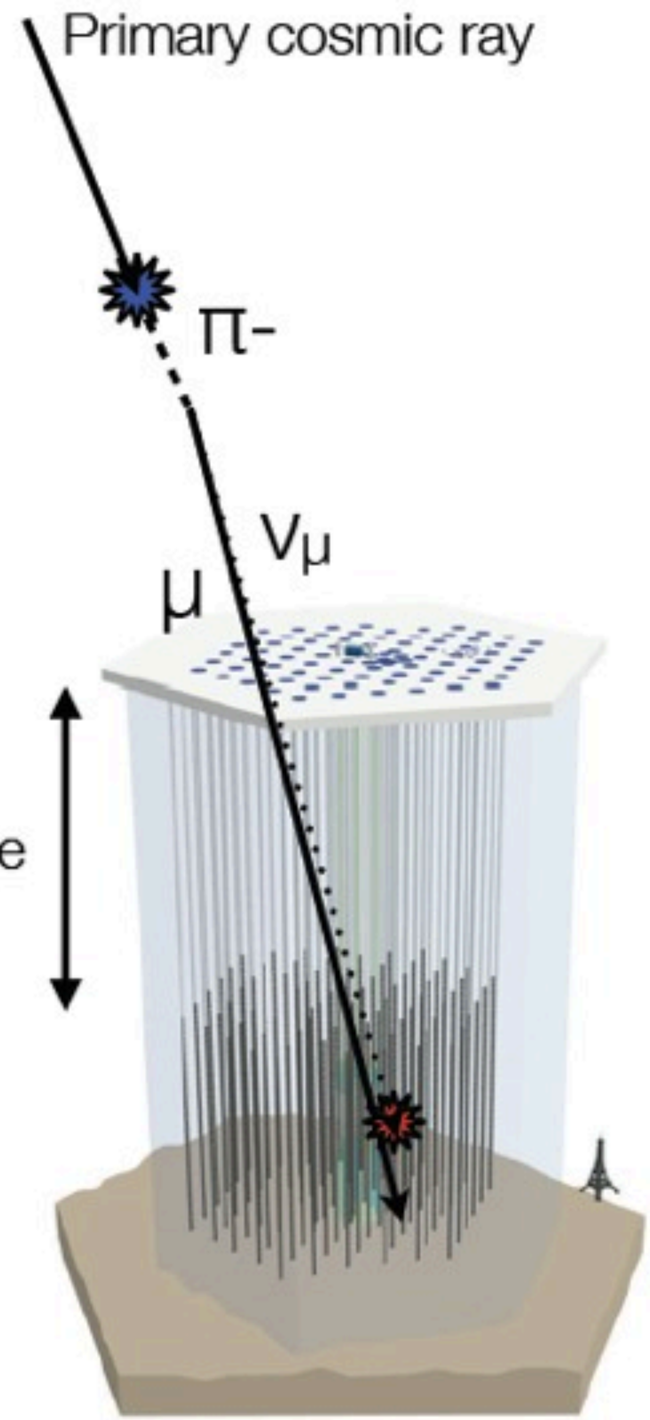
atmospheric neutrinos
down-ward veto (self-veto)

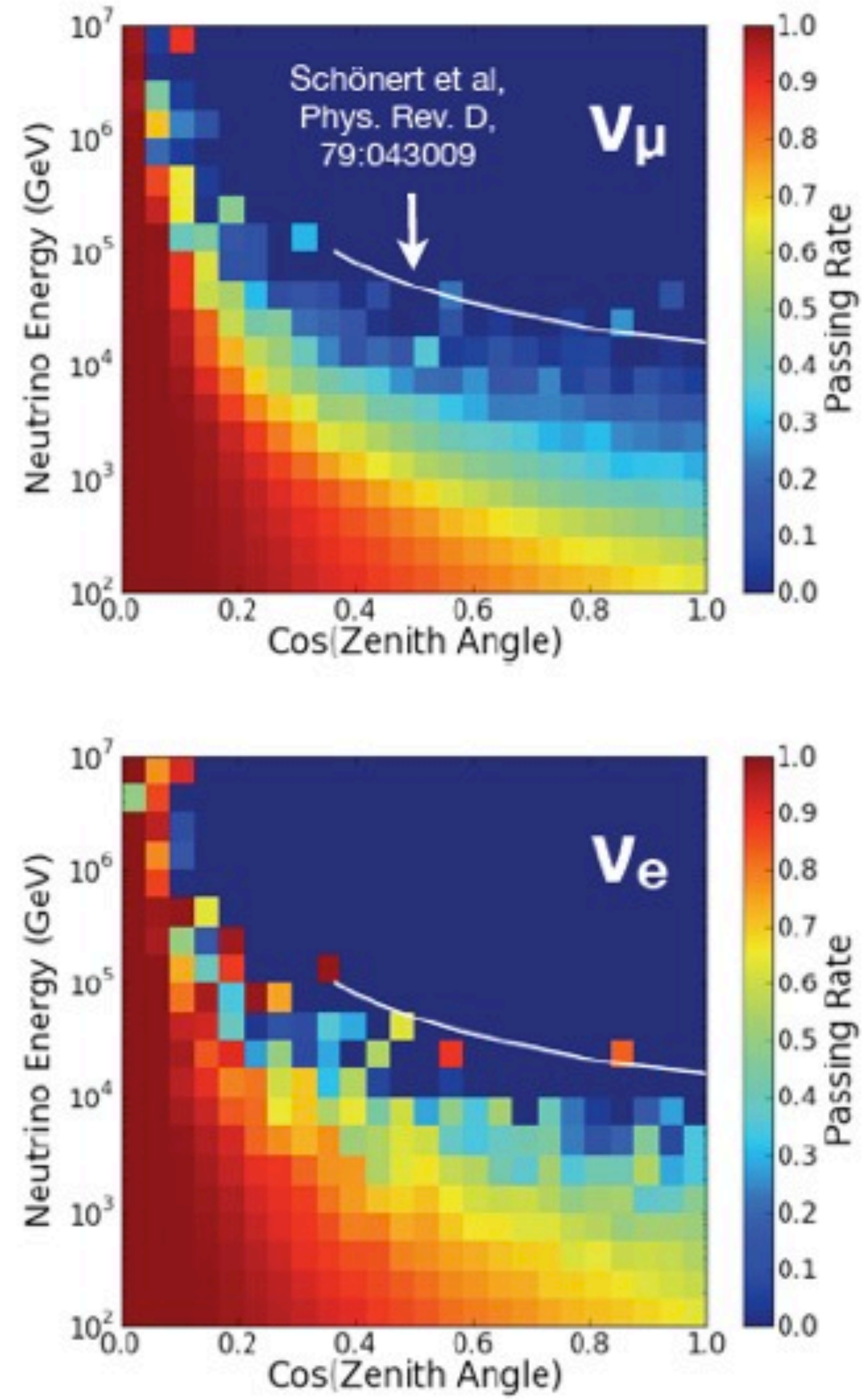
Jakob van Santen

Atmospheric muons and neutrinos are produced in the same processes.

Sufficiently vertical/high-energy atmospheric neutrinos come with accompanying muons!

The Schönert et al. calculation was based on **survival probability only.** What if the accompanying muon goes undetected?





Plots courtesy of K. Jero (UW-Madison)

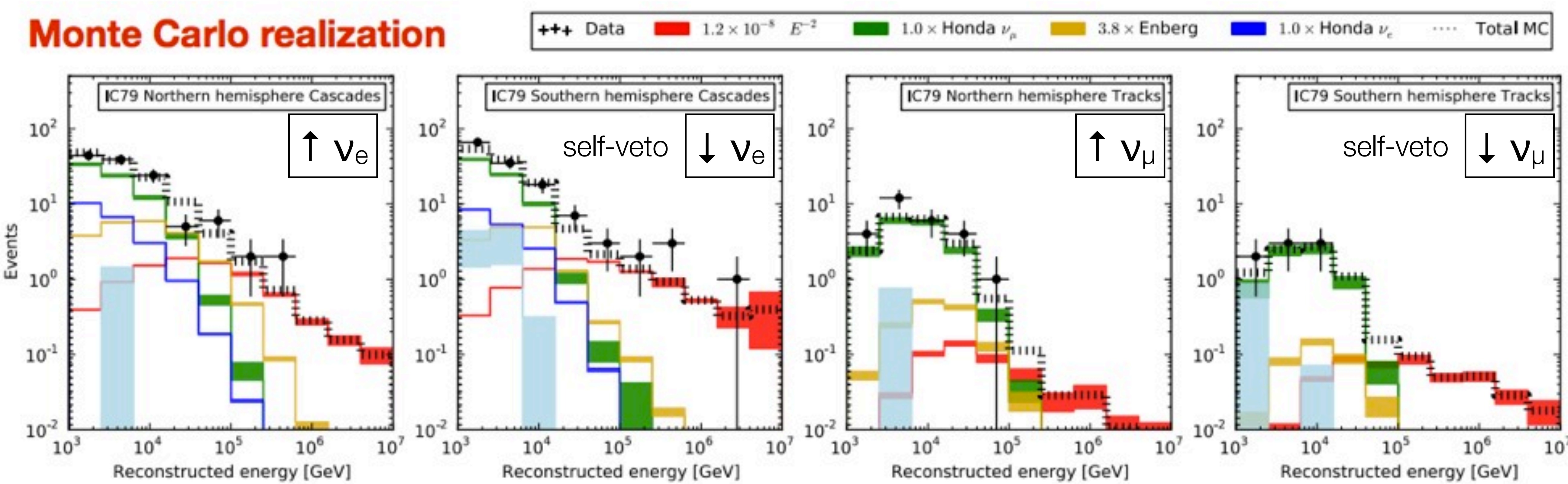
atmospheric neutrinos
charm and astrophysical neutrinos

- search for **high-energy all-flavor** neutrinos interacting inside (**starting**) the IceCube km3 instrumented volume from **all directions**

- charm & astrophysical signal are degenerate

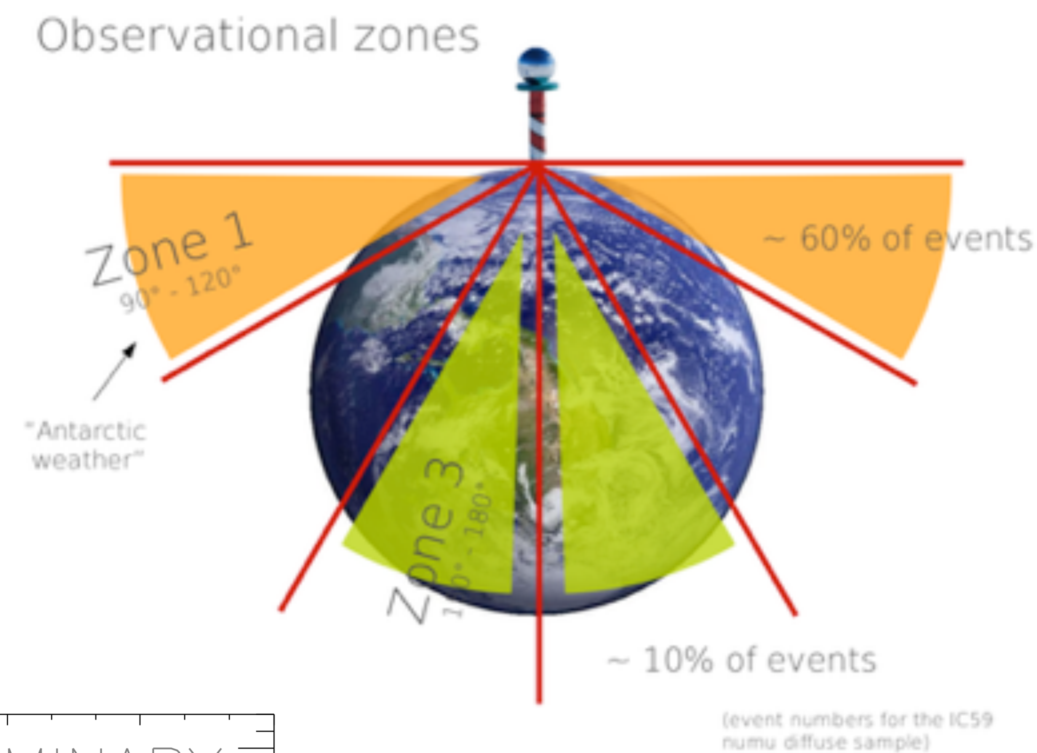
- need to determine charm contribution: multi-flavor global fit

J. van Santen - TAUP 13



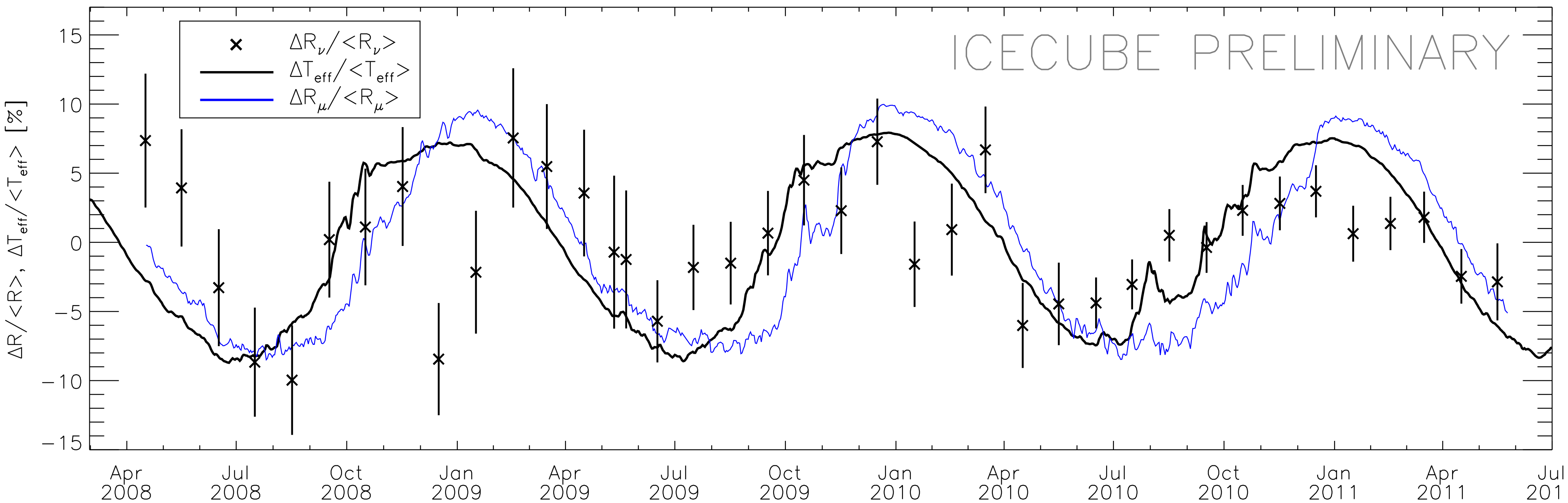
atmospheric neutrinos

charm & ν seasonal variations



effective temperature

$$T_{eff}(\theta) = \frac{\int dE_\nu \int dX\ A_{eff}(E_\nu, \theta) P(E_\nu, \theta, X) T(\theta, X)}{\int dE_\nu \int dX\ A_{eff}(E_\nu, \theta) P(E_\nu, \theta, X)}$$



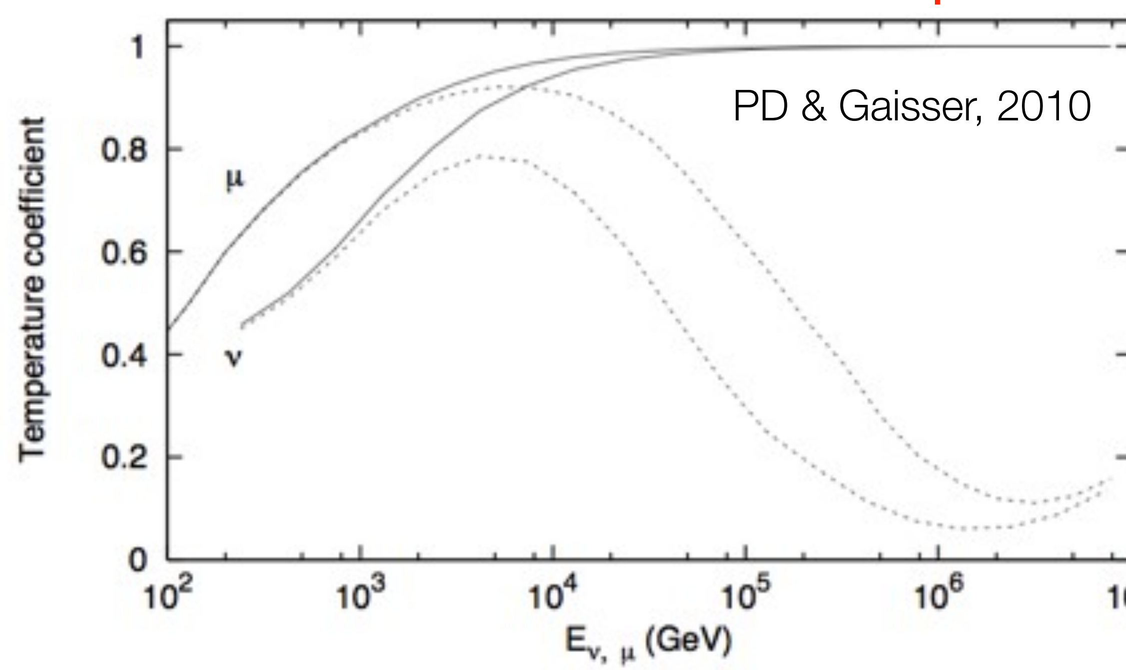
Tilav et al., ICRC 2009
PD et al., ICRC 2011
PD et al., ICRC 2013

seasonal variations decrease with prompt component

temperature coefficient

$$\alpha_T^{th}(\theta) = \frac{T \cdot \frac{\partial}{\partial T} \int dE_\nu\, \phi_\nu(E_\nu, \theta) A_{\rm eff}(E_\nu, \theta)}{\int dE_\nu\, \phi_\nu(E_\nu, \theta) A_{\rm eff}(E_\nu, \theta)}$$

$$\frac{\Delta R_\nu}{\langle R_\nu \rangle} = \alpha_T^{exp} \frac{\Delta T_{\rm eff}}{\langle T_{\rm eff} \rangle}$$

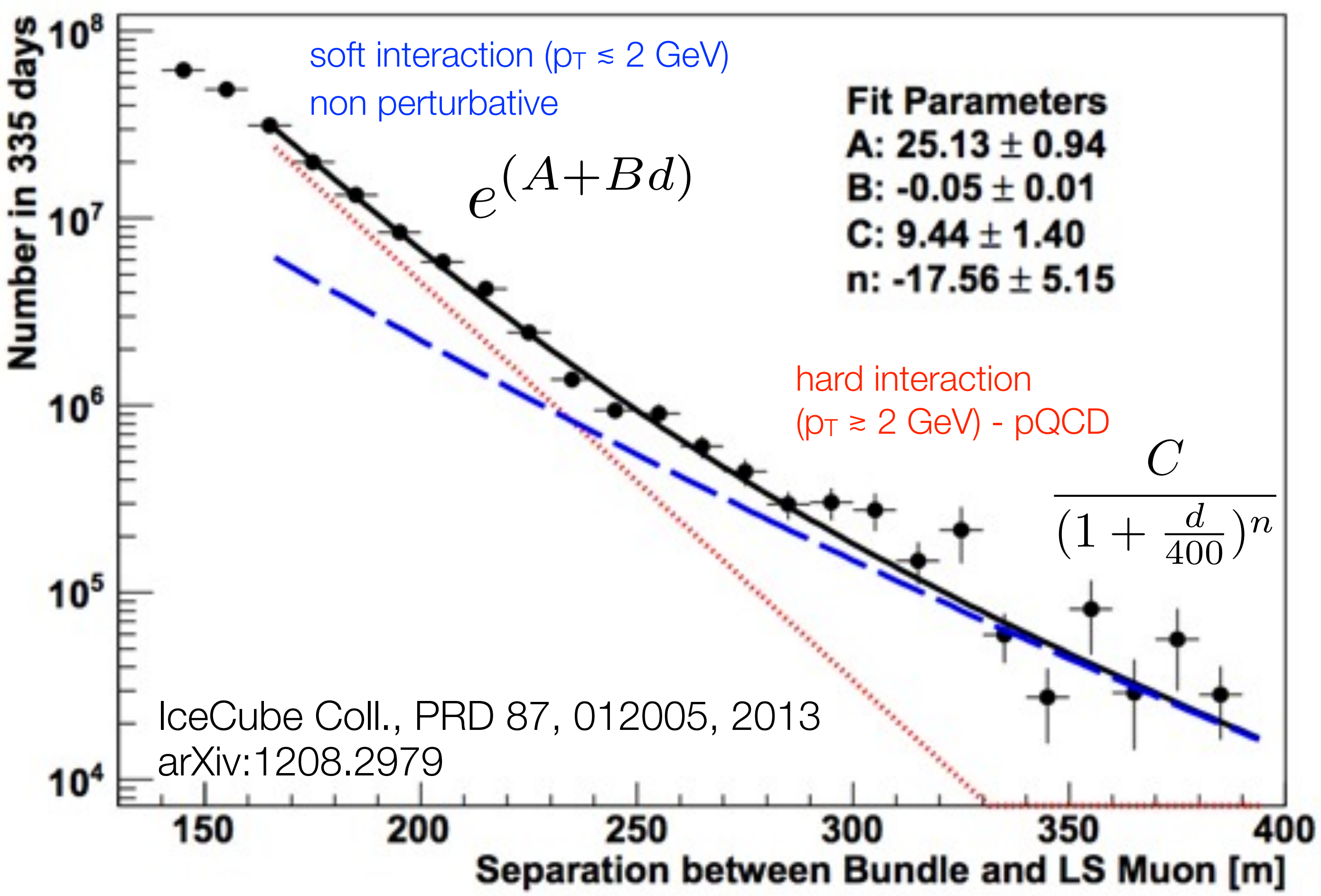


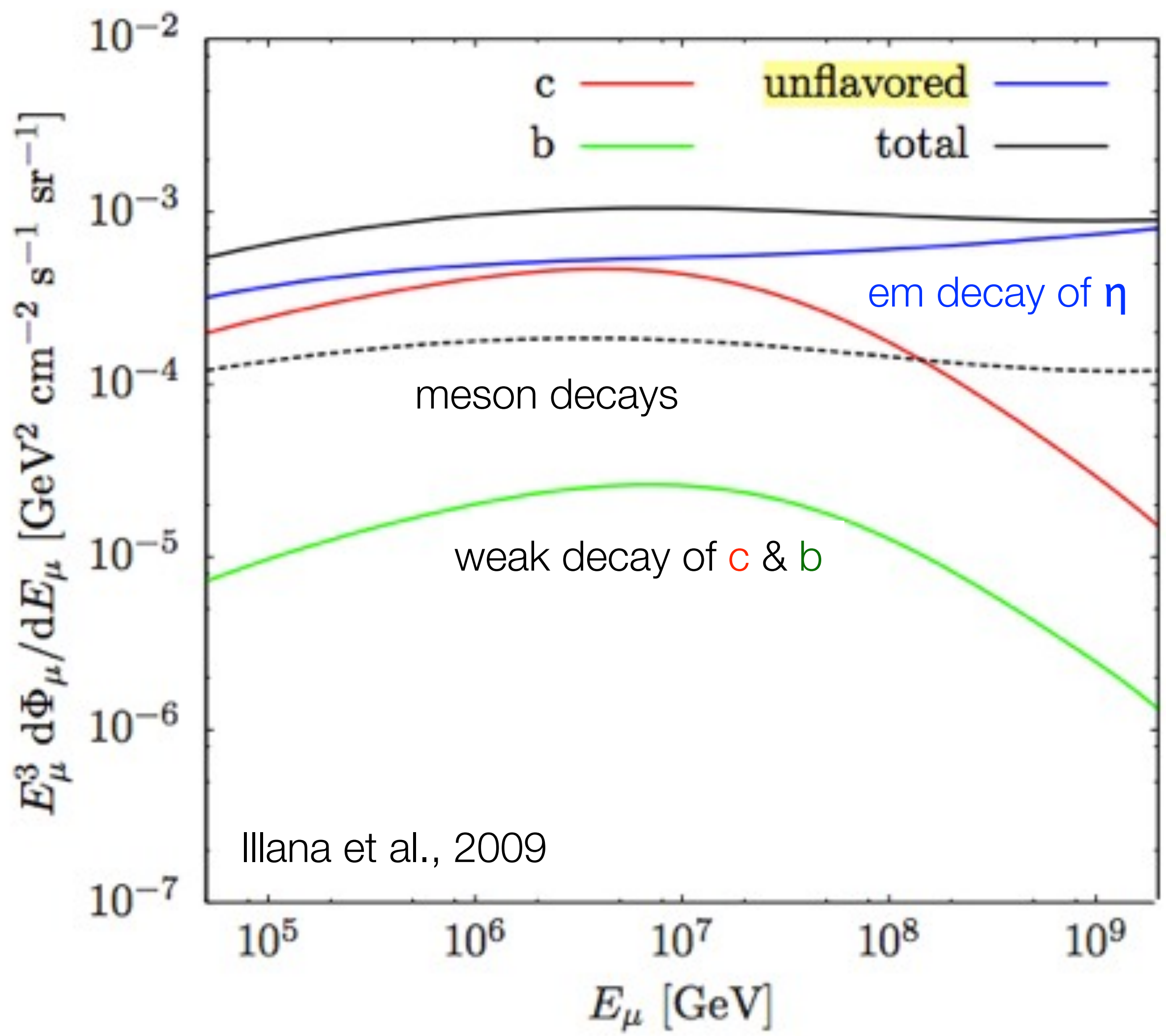
μ not *contaminated* by astrophysical signal BUT multiplicity

charm production in the atmosphere

breaking the degeneracy with astrophysical signal

- μ produced by **same processes** as ν
- BUT not *contaminated* by astro signal
- hadronic model **uncertainties** (effect of unflavoured mesons)

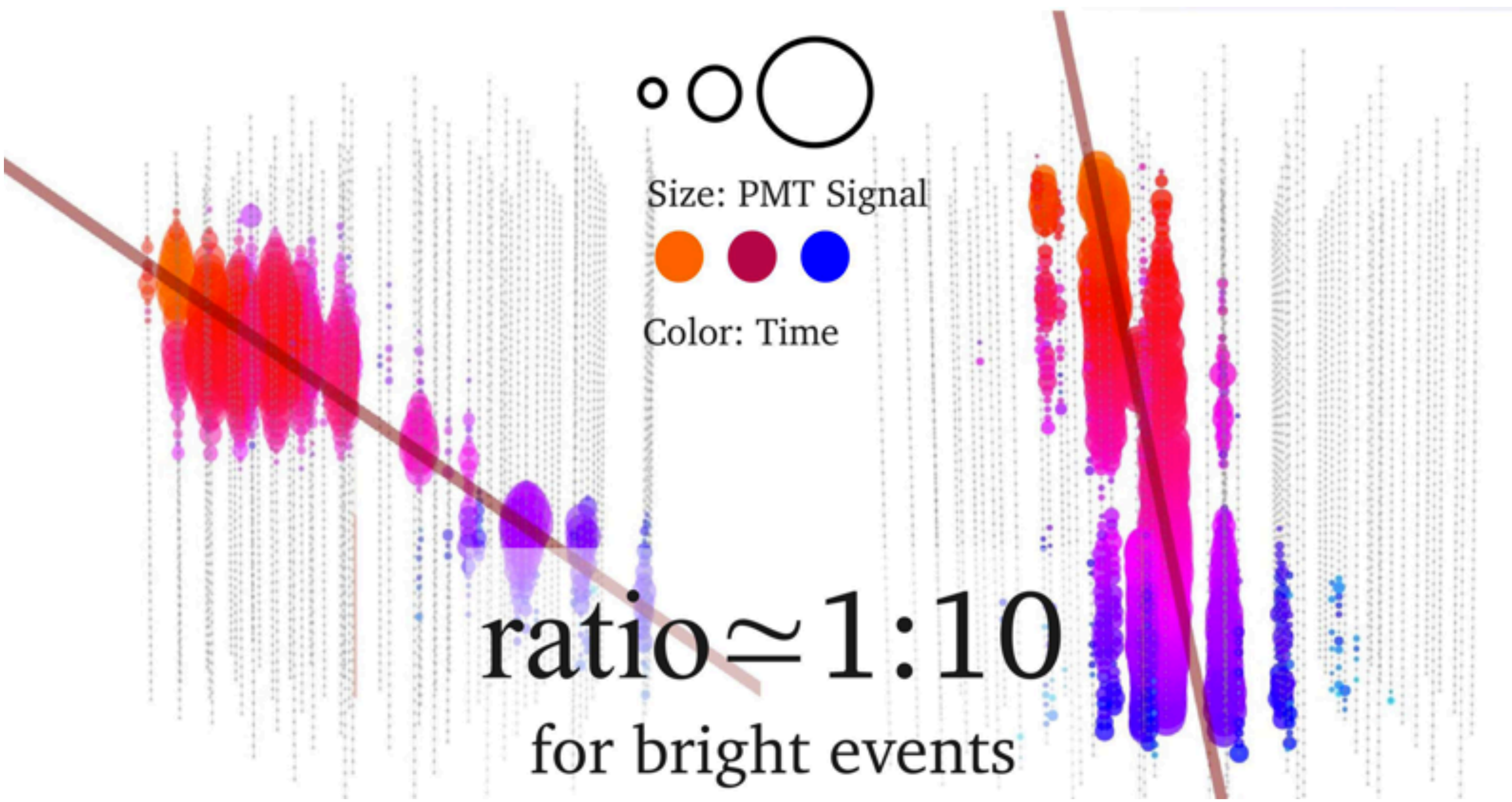




$$d_T \approx \frac{p_T H c}{E_\mu cos(\theta)}$$

charm production in the atmosphere
breaking the degeneracy with astrophysical signal

- separate μ bundles (smooth) from high energy μ (stochastics)
- measure energy spectrum up to PeV scale



HE muon   High multiplicity bundle

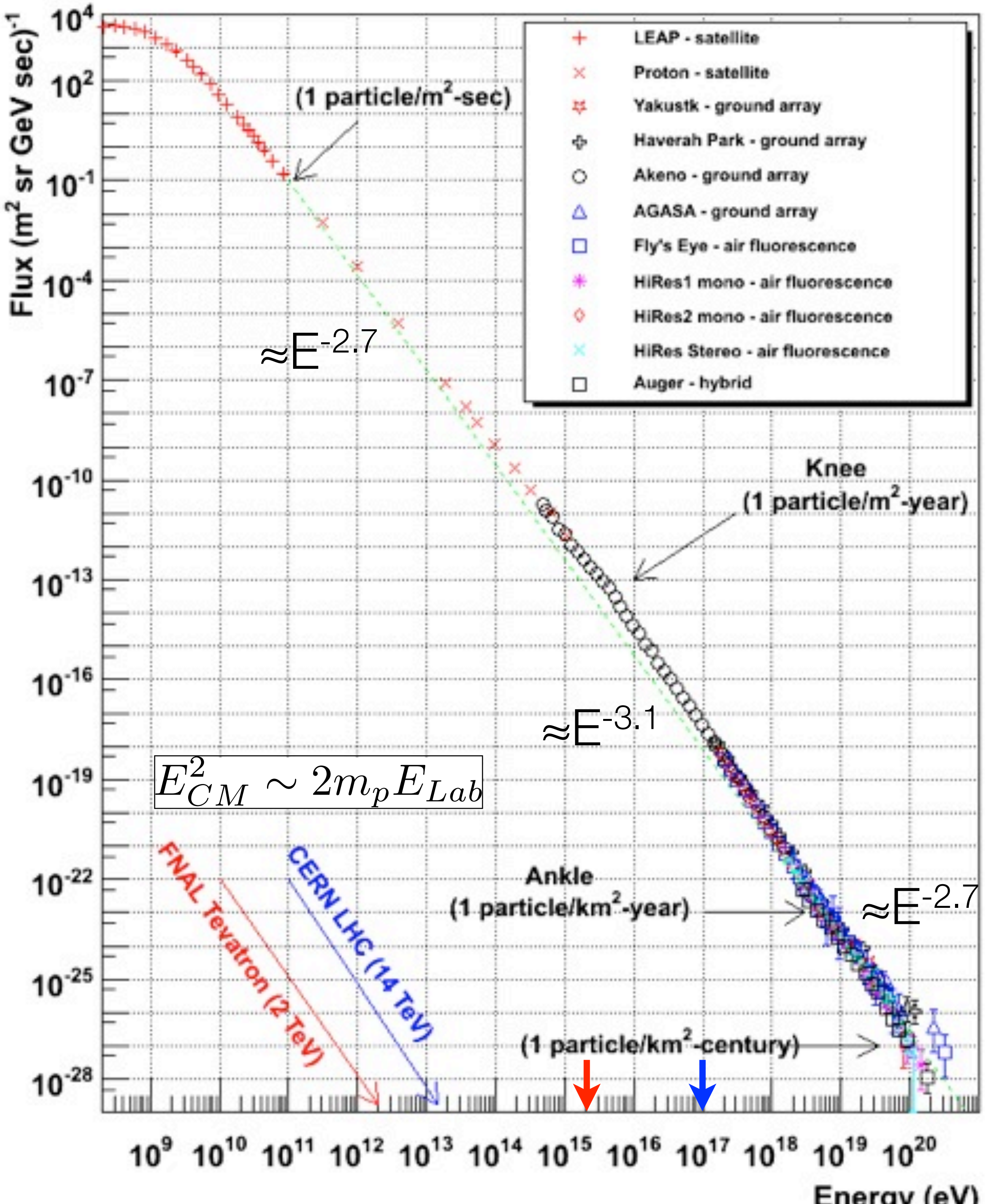
Paolo Desiati

simulating CR-induced background

- CORSIKA to generate atmospheric muons and neutrinos
- 5 mass groups: re-weight with arbitrary CR composition
- hadronic interaction models with charm production: recent updates from LHC results
  - consistent treatment of muons and neutrinos (for prompt measurements)
  - correlated systematics in reconstructed muon / neutrino events
  - correct general treatment of self-veto for any analysis (different thresholds)
    - 3rd party generated neutrinos input to NeutrinoGenerator
    - track parent particle ID
    - weight meson decay for higher efficiency

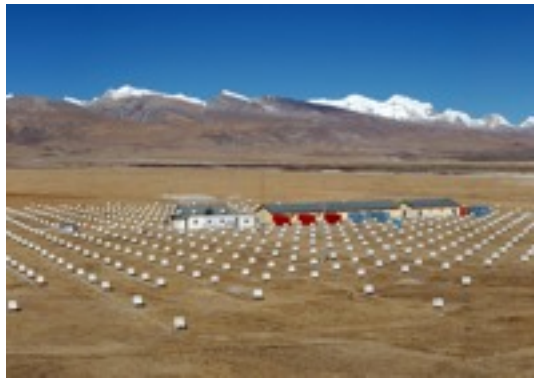
backup slides

primary cosmic rays
spectrum & composition

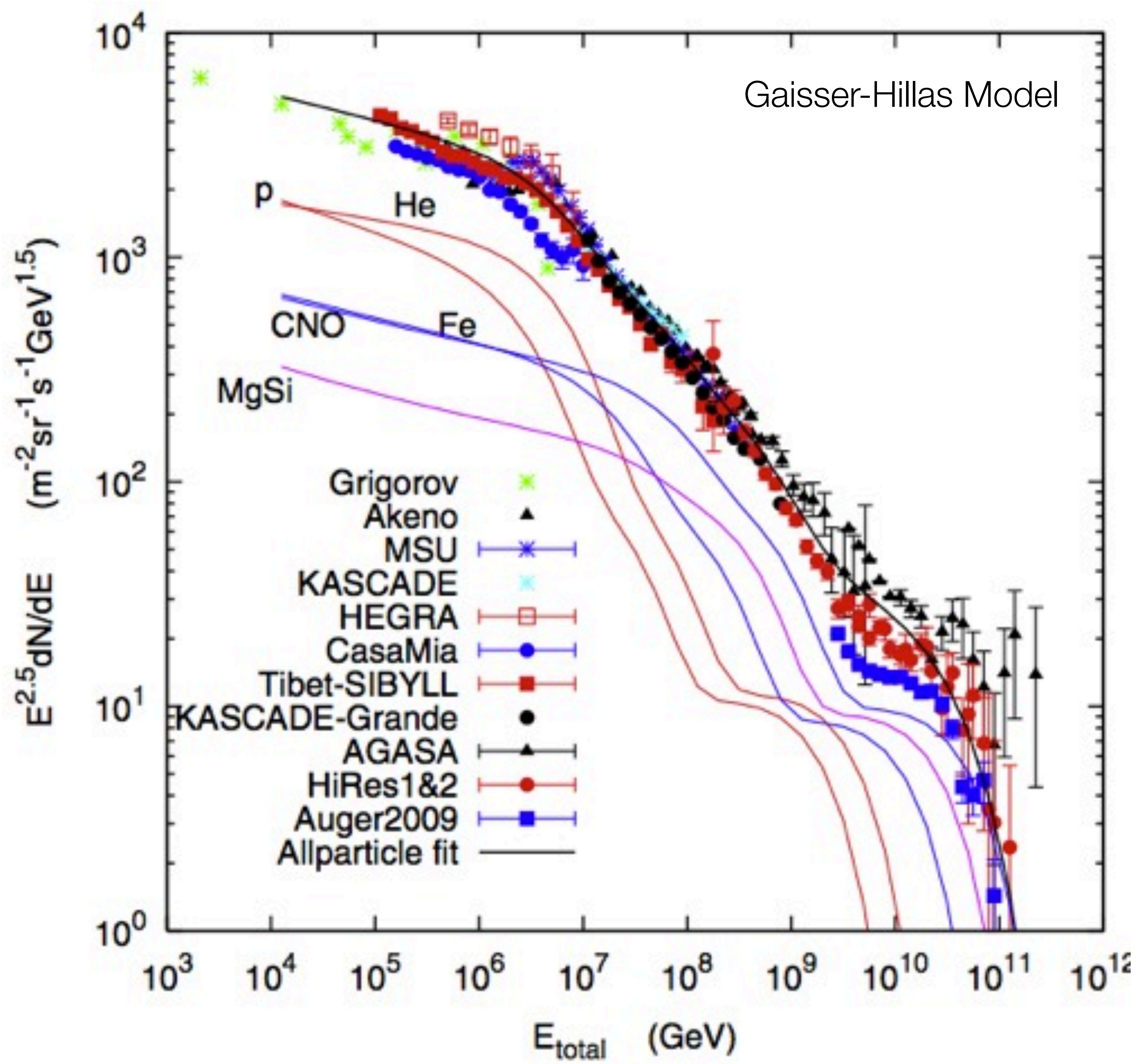


direct measurements

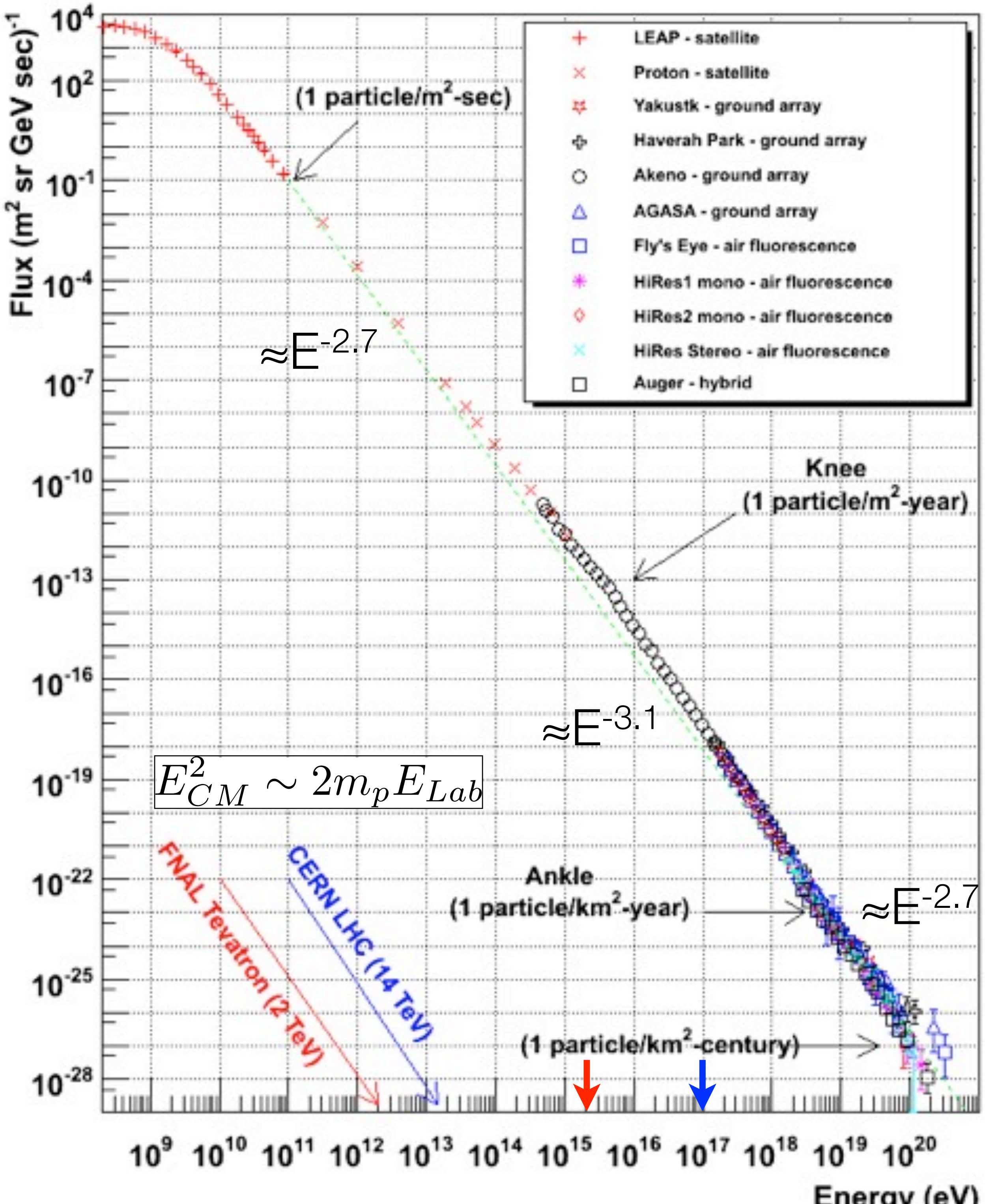
indirect measurements



Gaisser, Astropart. Phys. 35 (2012) 801

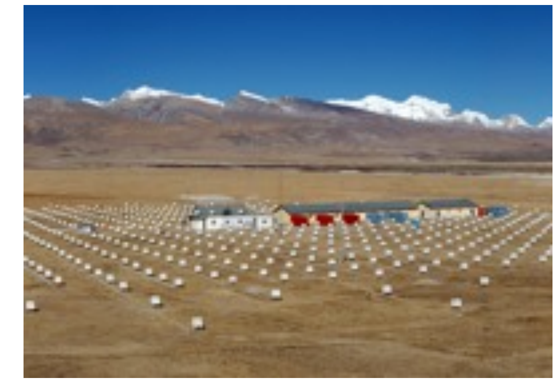


primary cosmic rays
spectrum & composition

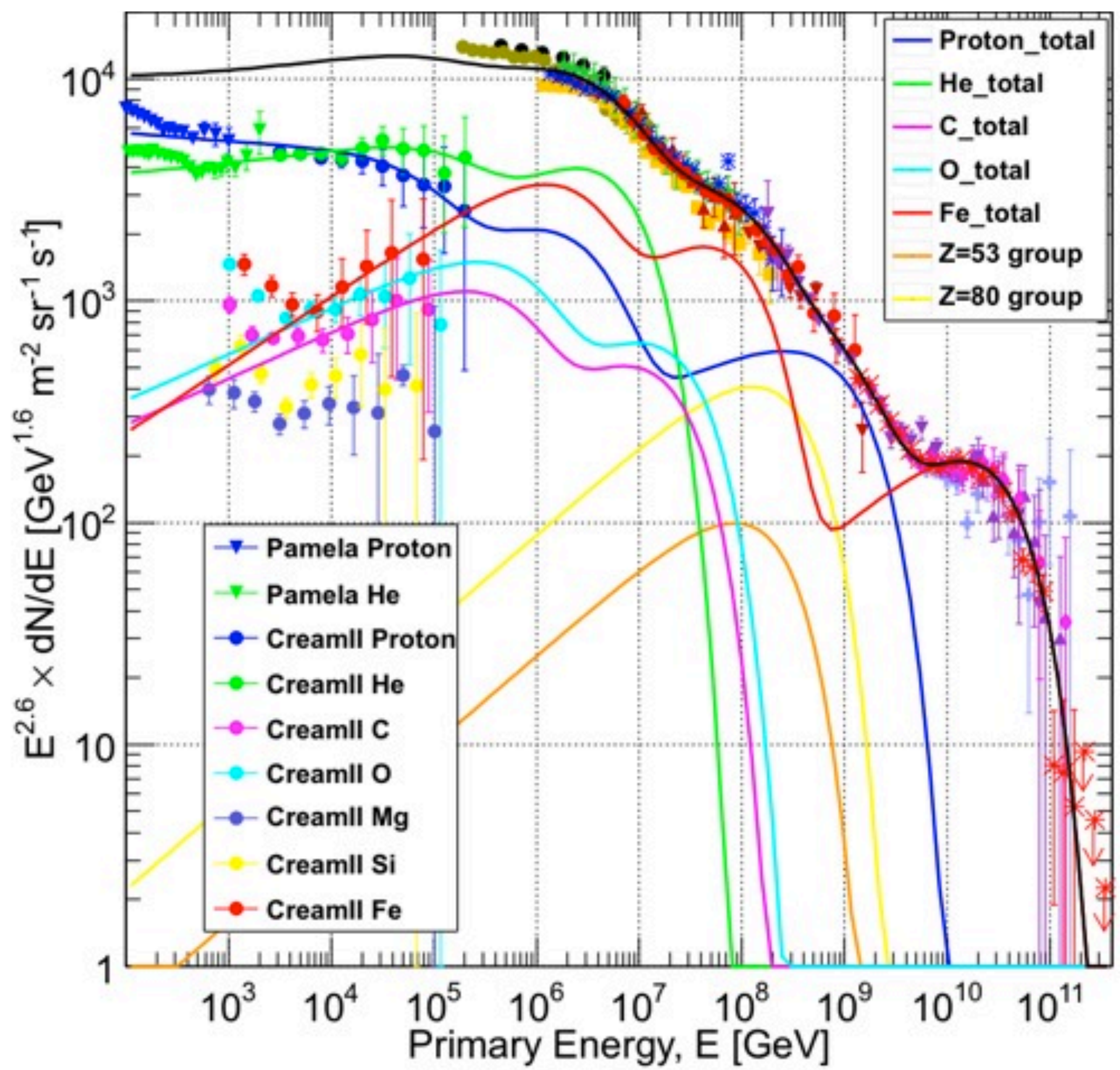


direct
measurements

indirect
measurements



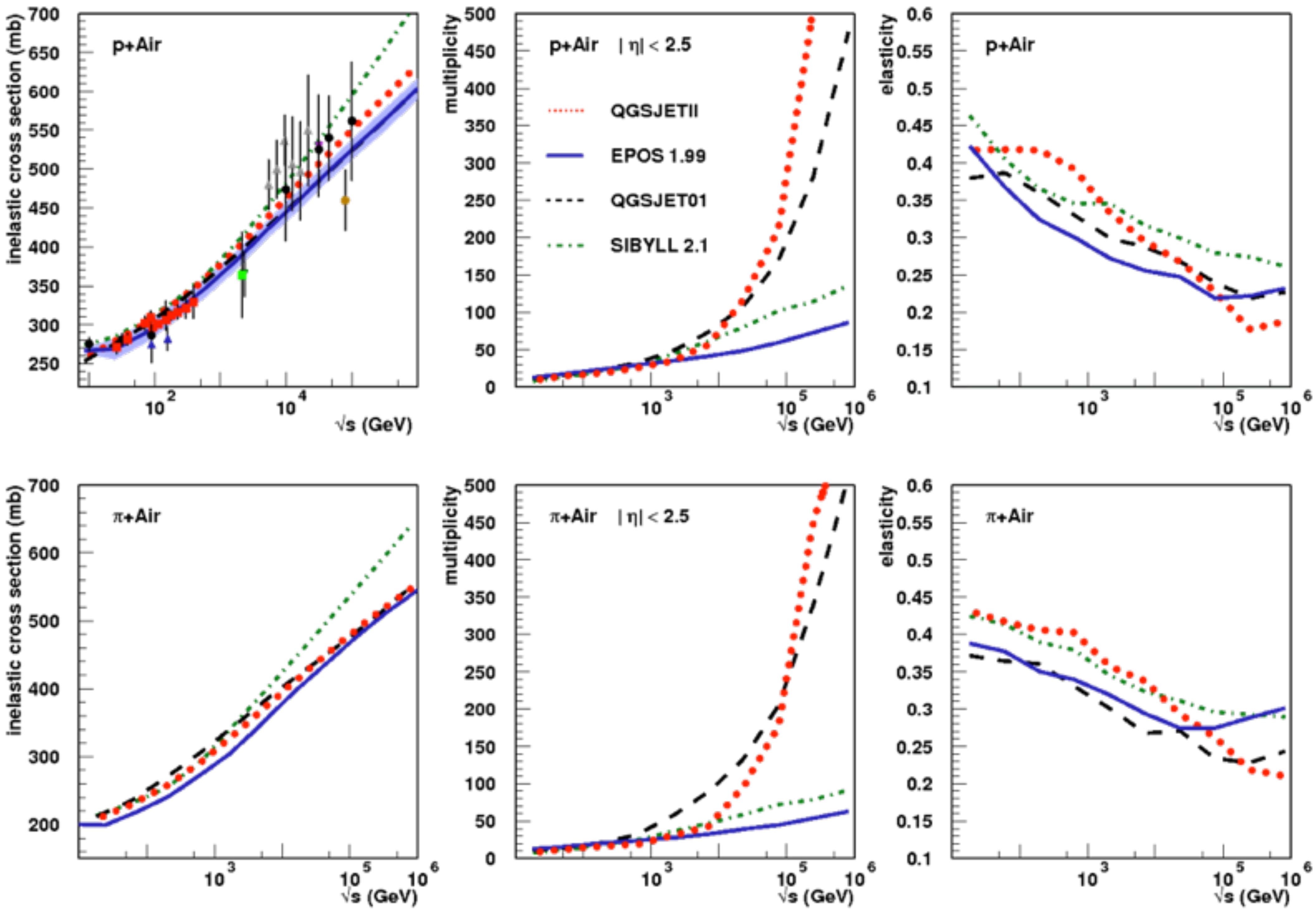
Gaisser, Stanev & Tilav, 2013 - arXiv:1303.3565



Paolo Desiati

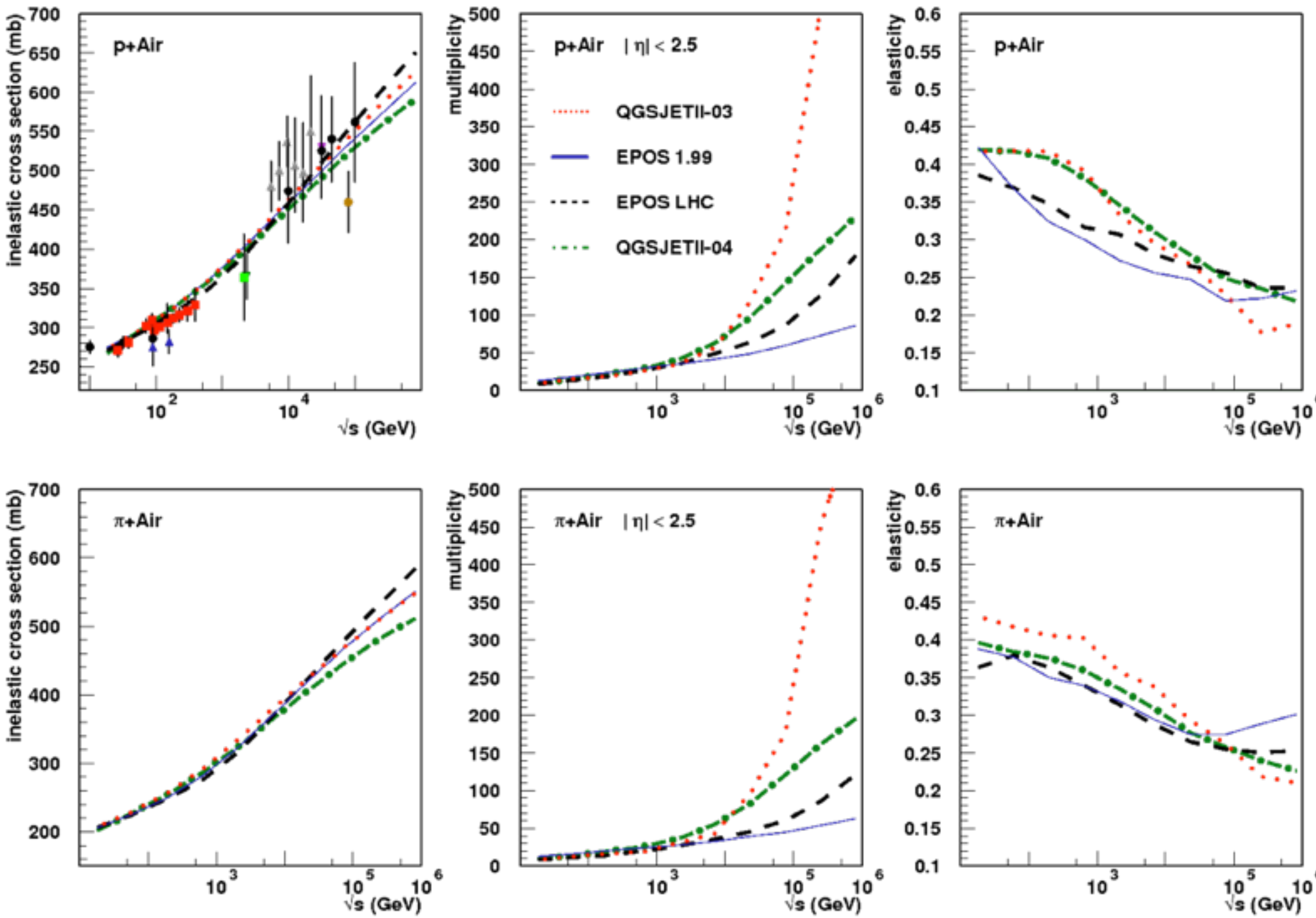
hadronic interactions

reduction of systematic uncertainties



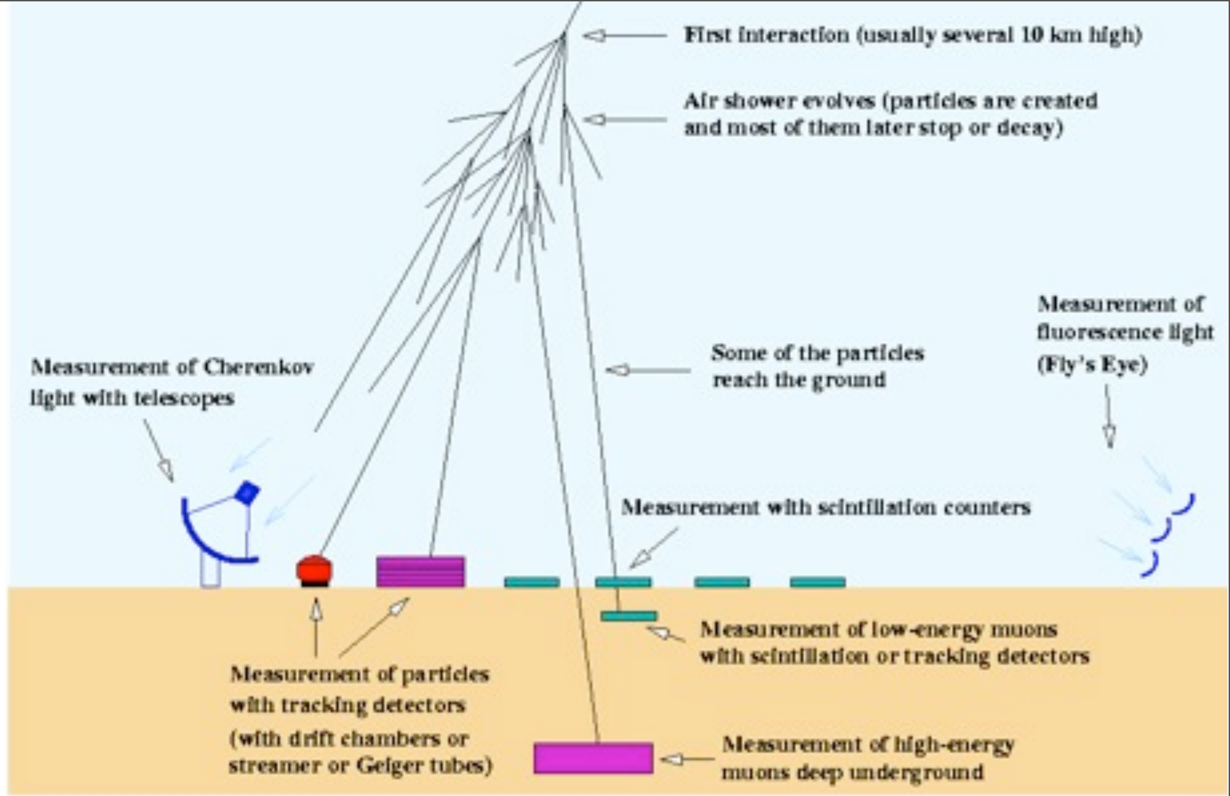
Paolo Desiati

hadronic interactions
reduction of systematic uncertainties

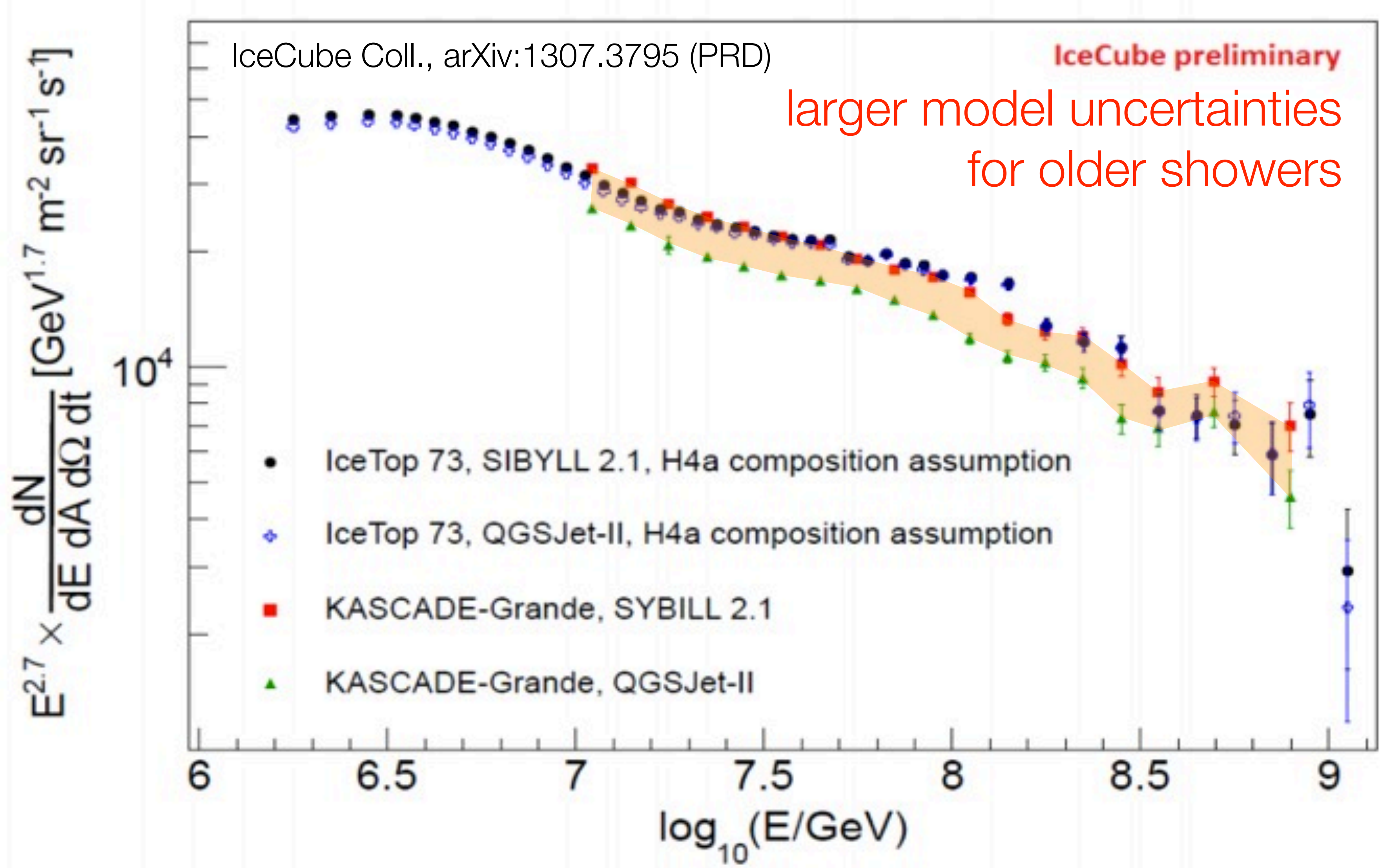


> 100s TeV cosmic rays

indirect observations



- e.m. & hadronic shower components observed at the Earth's surface
- measure energy deposited, temporal, longitudinal & lateral distributions, and unfold the primary energy & mass
- *KASCADE @ sea level*
- *IceTop @ 2800 m asl*



Monday, October 14, 2013

> 100 PeV cosmic rays

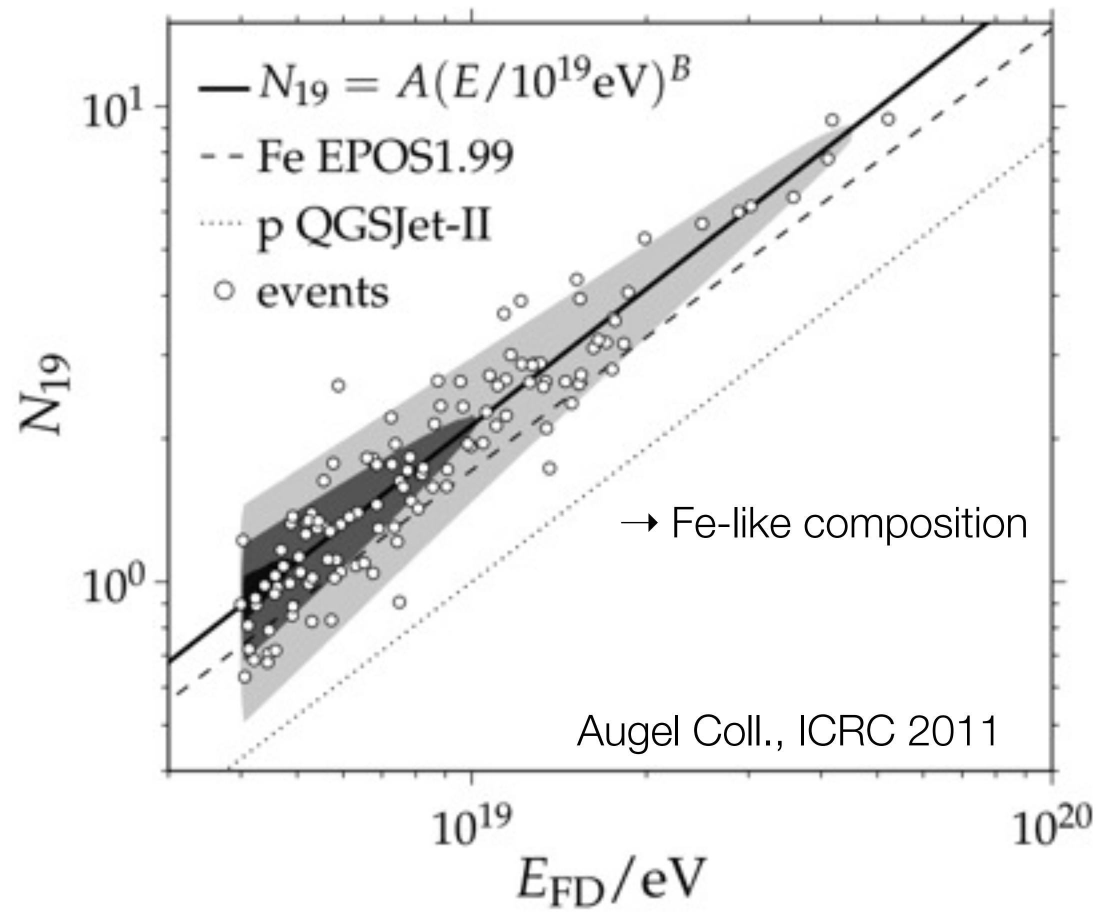
- inclined showers develop earlier and exhaust higher in the atmosphere

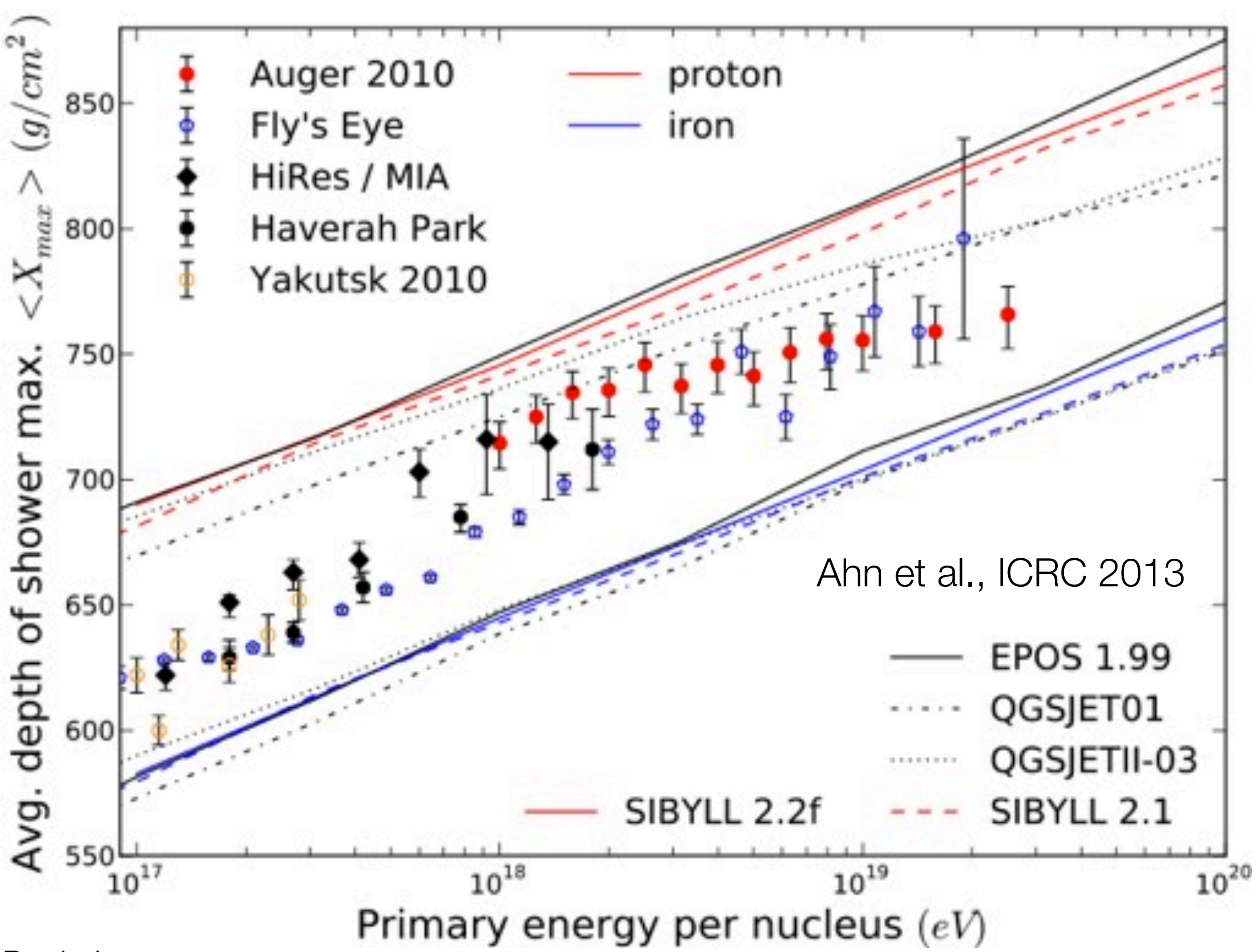
- only penetrating muons reach the ground

- **higher μ flux** observed above 10^{18} eV

  - N_{19}/QGSJet-II(10^{19} eV) = **2.13**±0.04±0.11 (sys.)

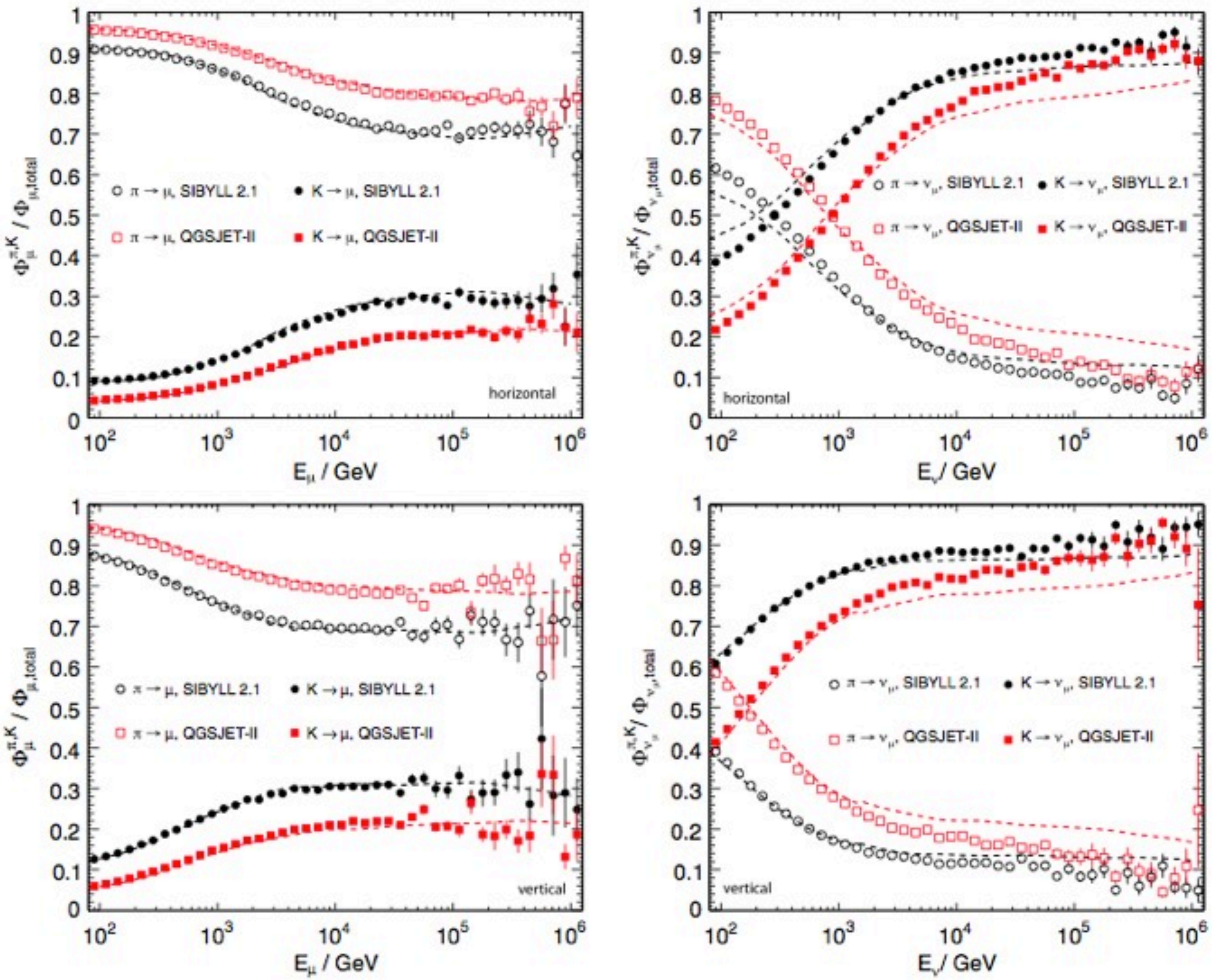
- **mass composition** affected by the large systematic uncertainties of interaction models (+ experimental techniques)





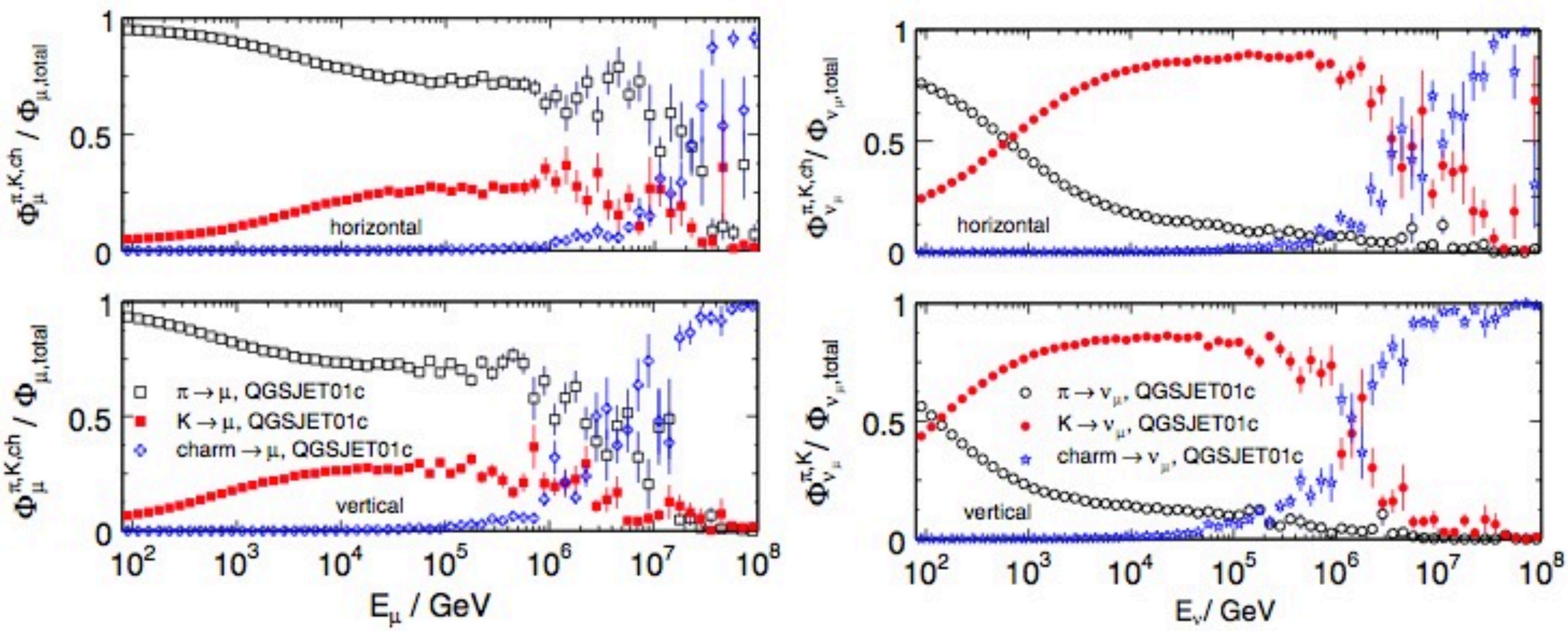
Fedynitch, Becker Tjus, PD 2012

atmospheric neutrinos
high energy and heavy quarks



atmospheric neutrinos
high energy and heavy quarks

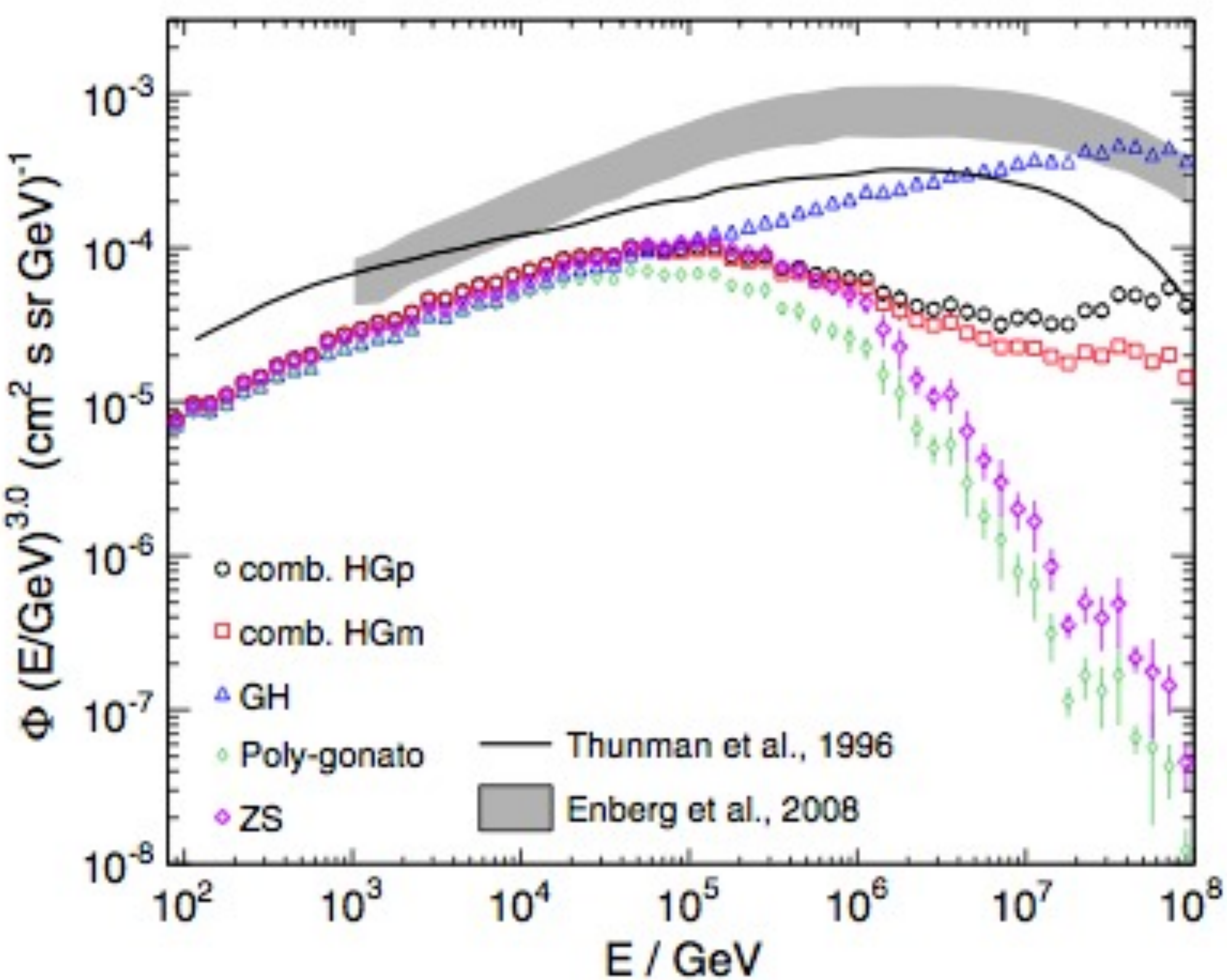
QGSJET01-c



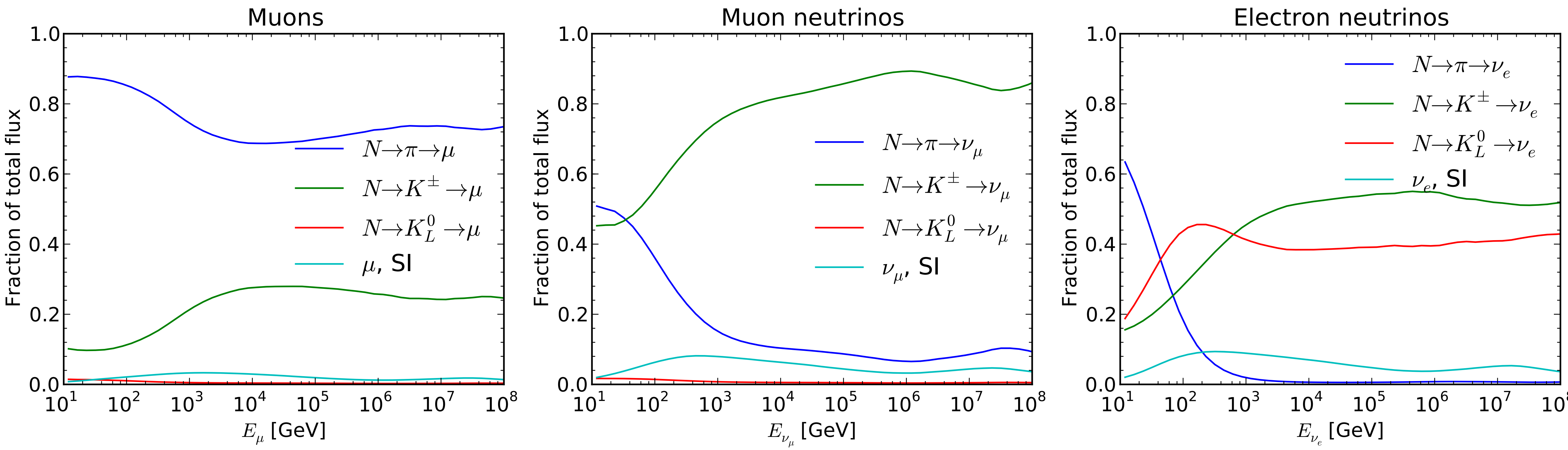
atmospheric neutrinos
high energy and heavy quarks

Fedynitch, Becker Tjus, PD 2012

QGSJET01-c



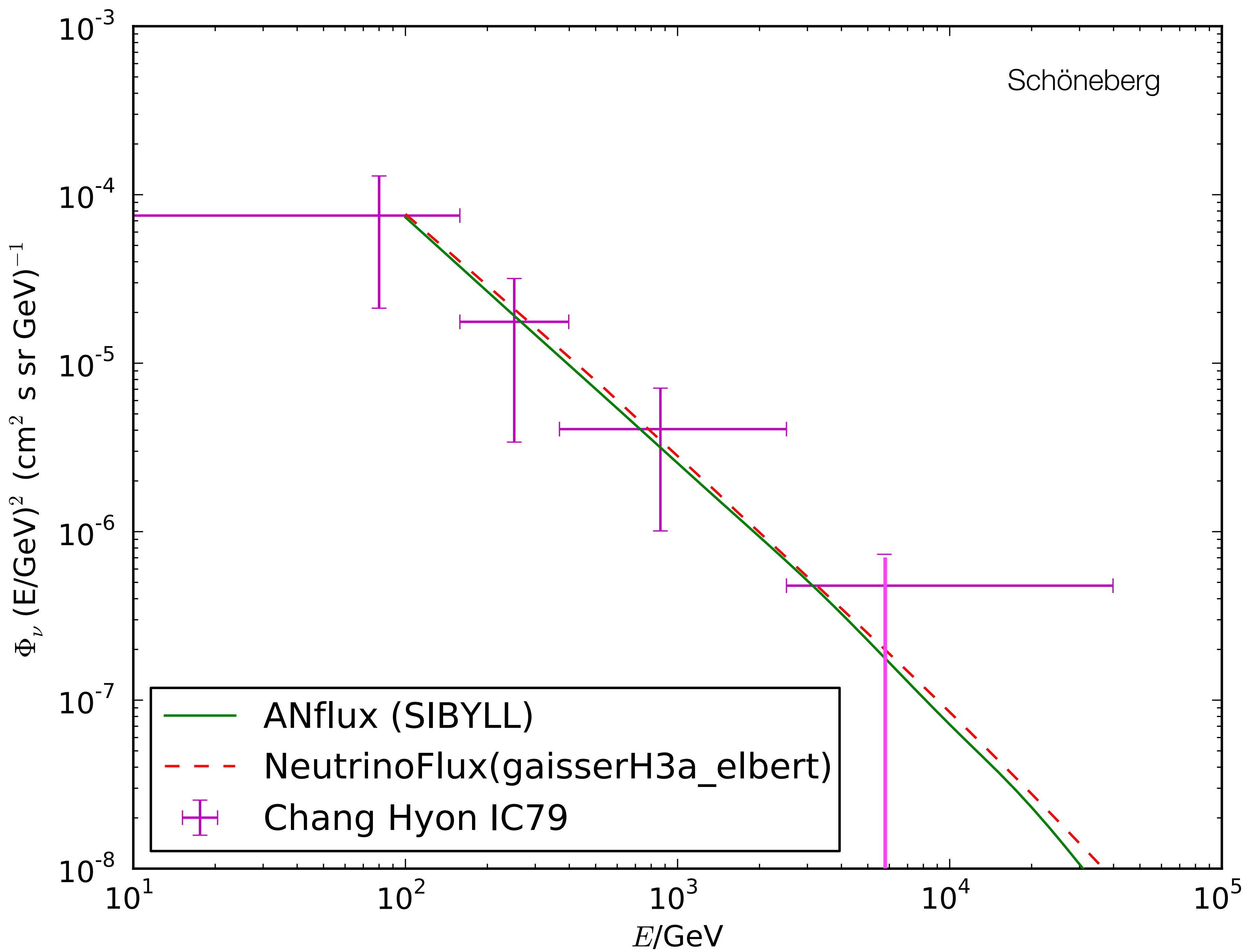
atmospheric neutrinos
high energy and heavy quarks



atmospheric neutrinos

current status

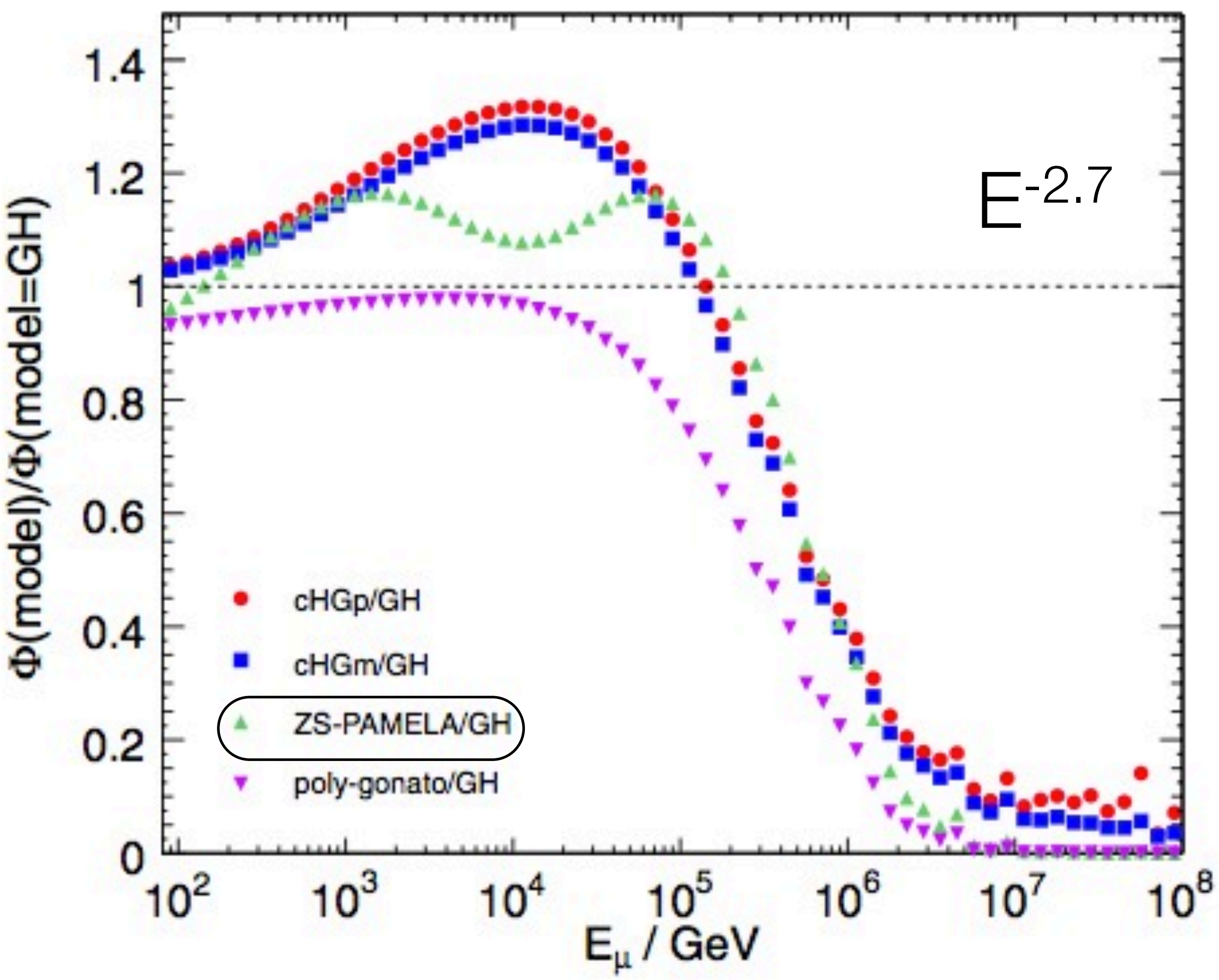
observed cascading $\nu_e + \bar{\nu}_e$

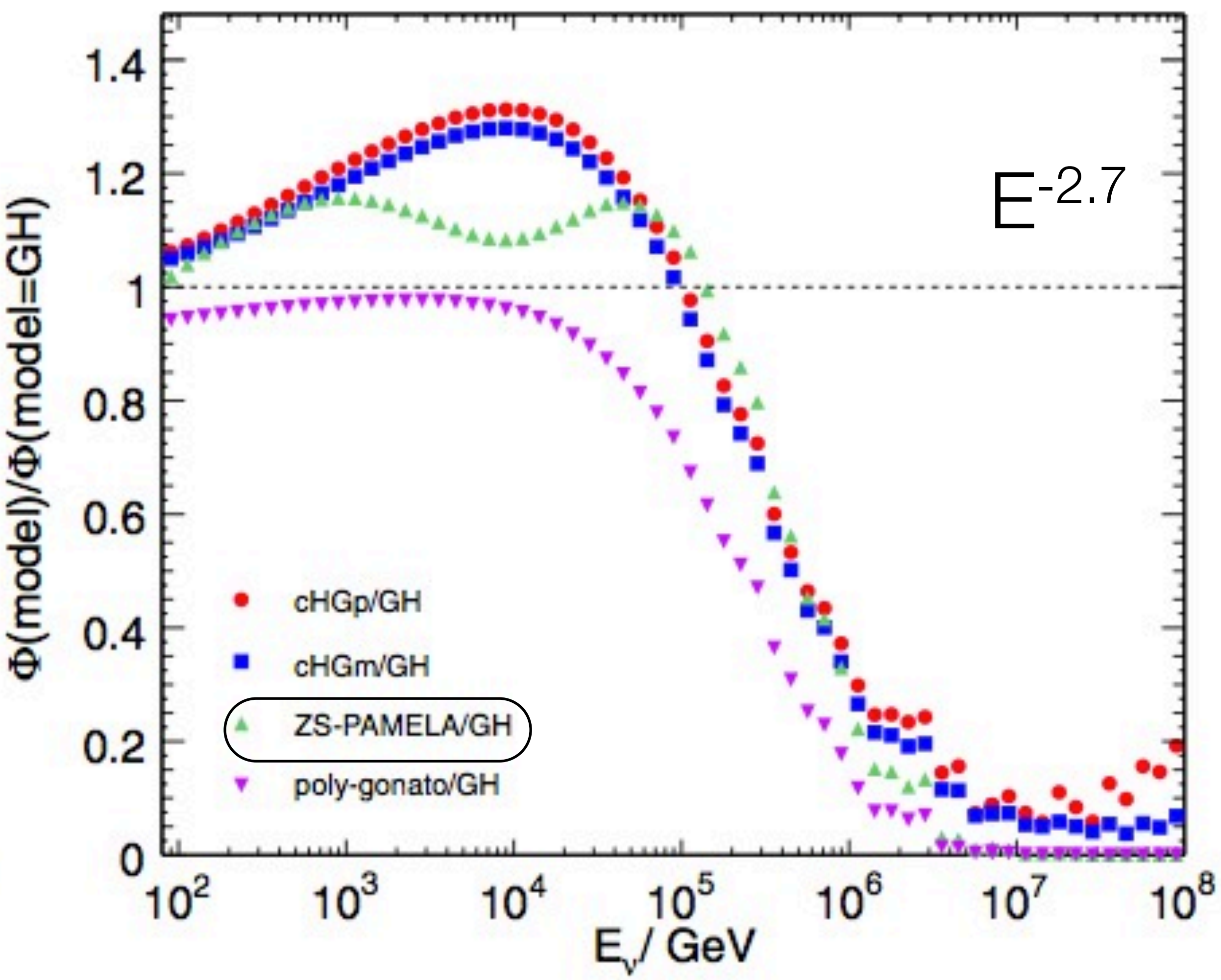


atmospheric muons and neutrinos

effect of cosmic ray spectrum

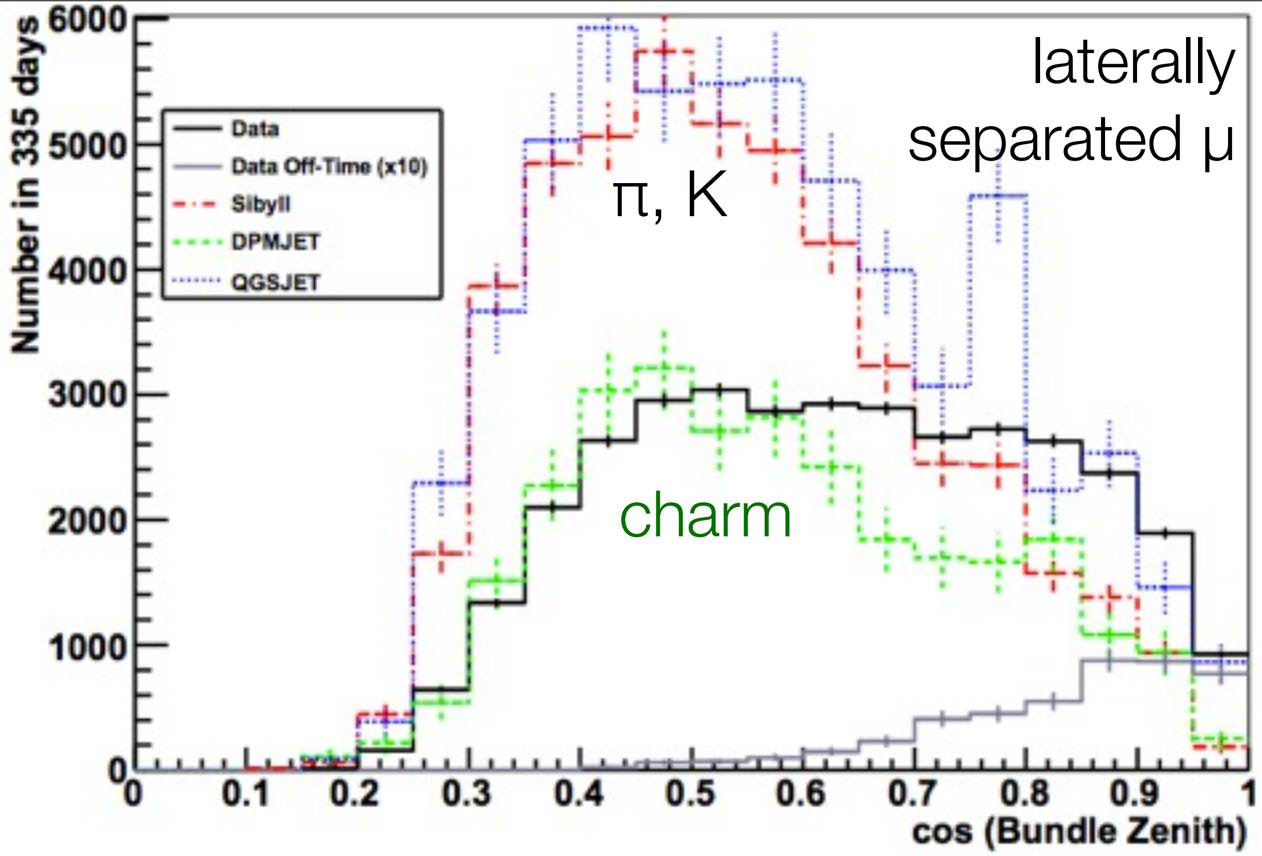
Fedynitch, Becker Tjus, PD 2012



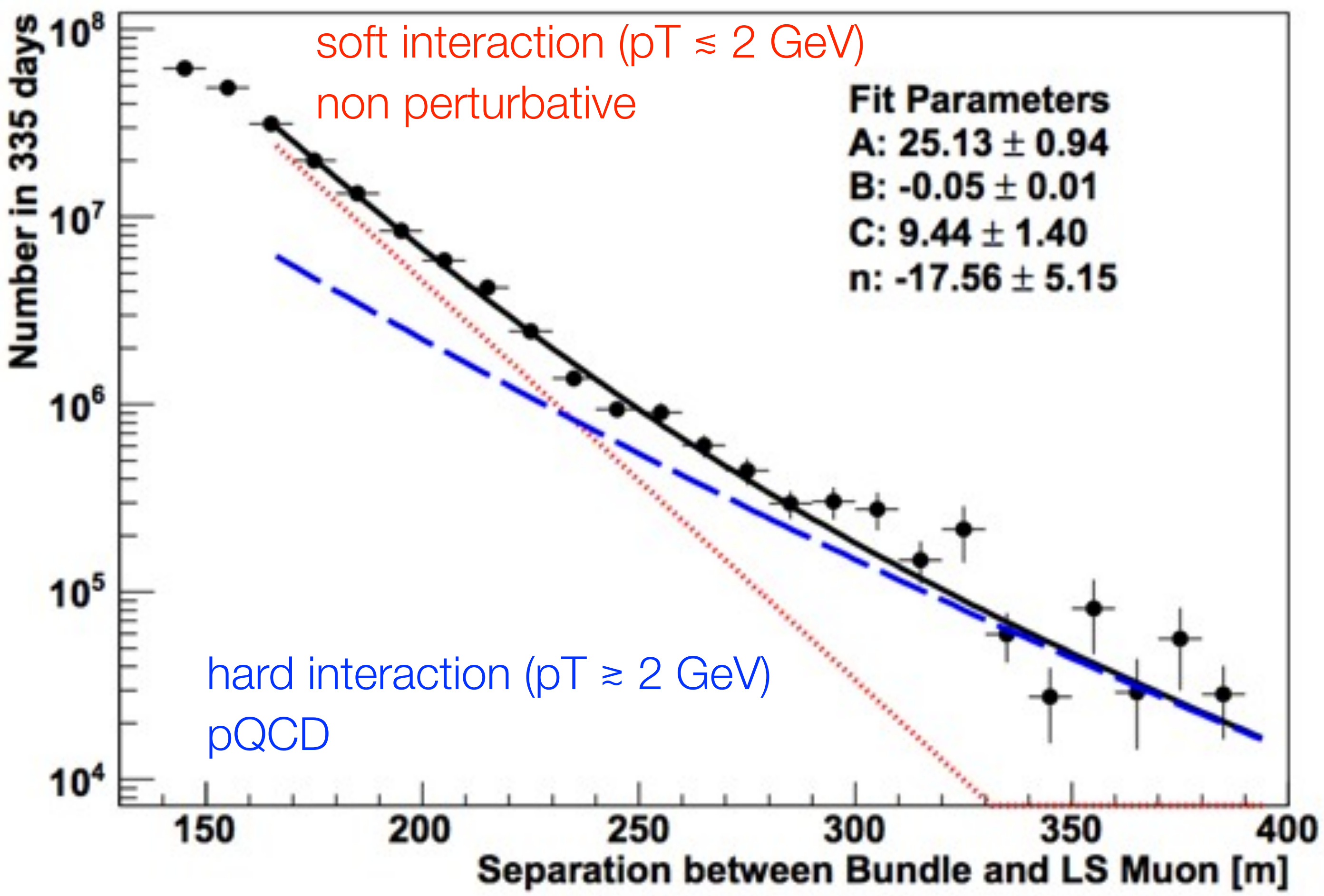


atmospheric neutrinos
charm and high p_T muons

- search for μ + μ bundle
- measure separation $$d_T \approx \frac{p_T H c}{E_\mu cos(\theta)}$$
- CR composition & interaction models



IceCube Coll., PRD 87, 012005, 2013
arXiv:1208.2979



- increased K and charm contribution
  - improve forward region
- lighter cosmic ray composition

siati

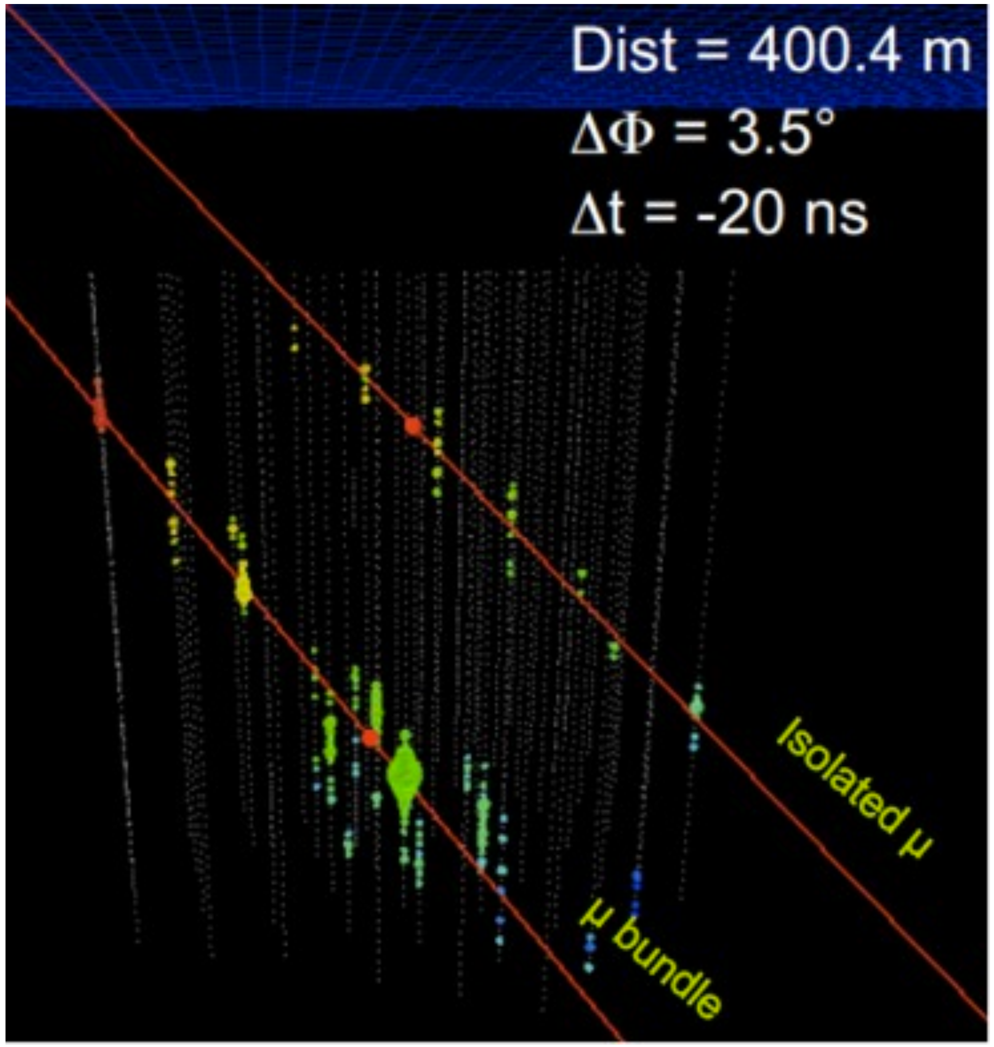
atmospheric neutrinos

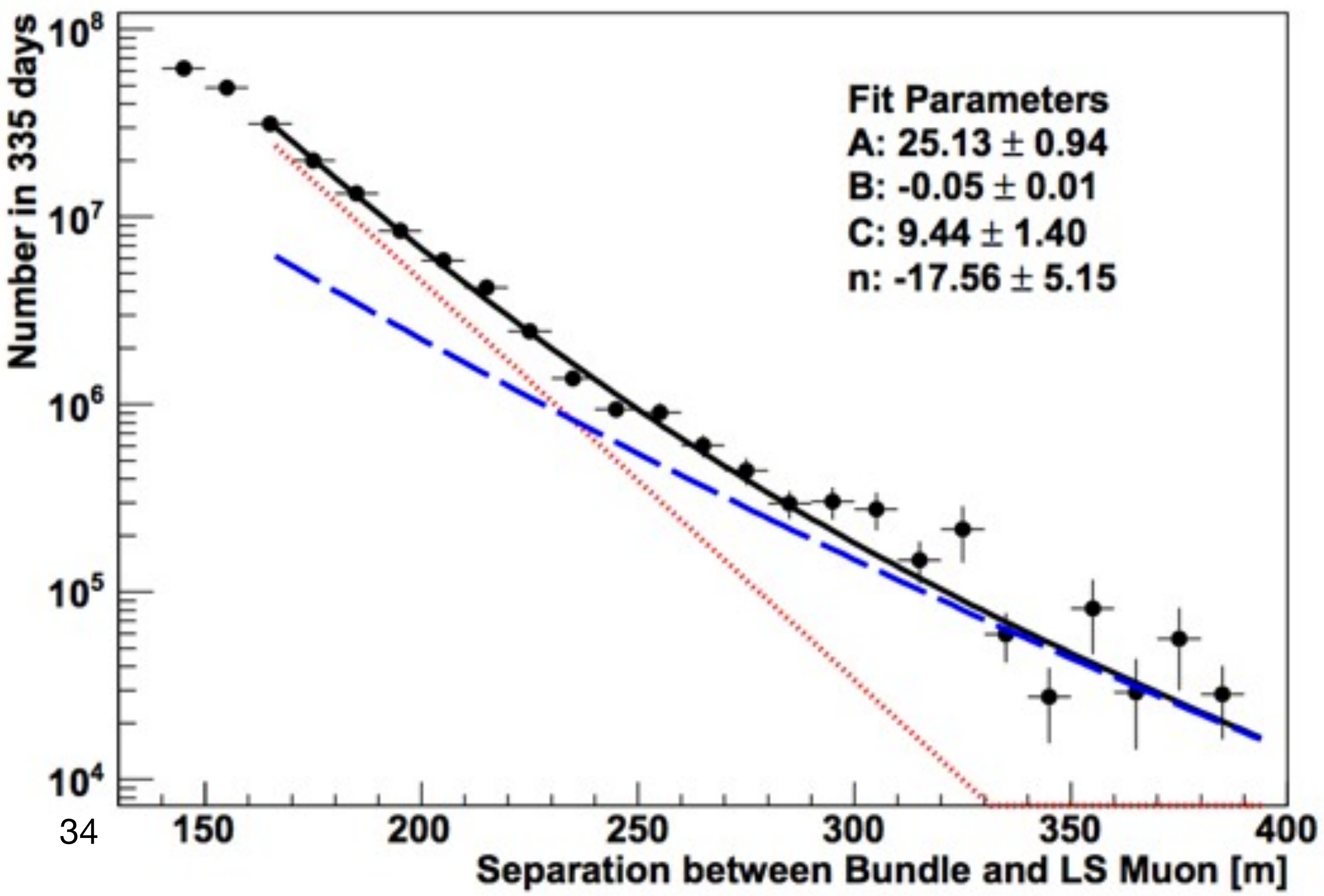
charm and high p_T muons

- search for μ + μ bundle
- measure separation

$$d_T \approx \frac{p_T H c}{E_\mu cos(\theta)}$$

- CR composition & interaction models





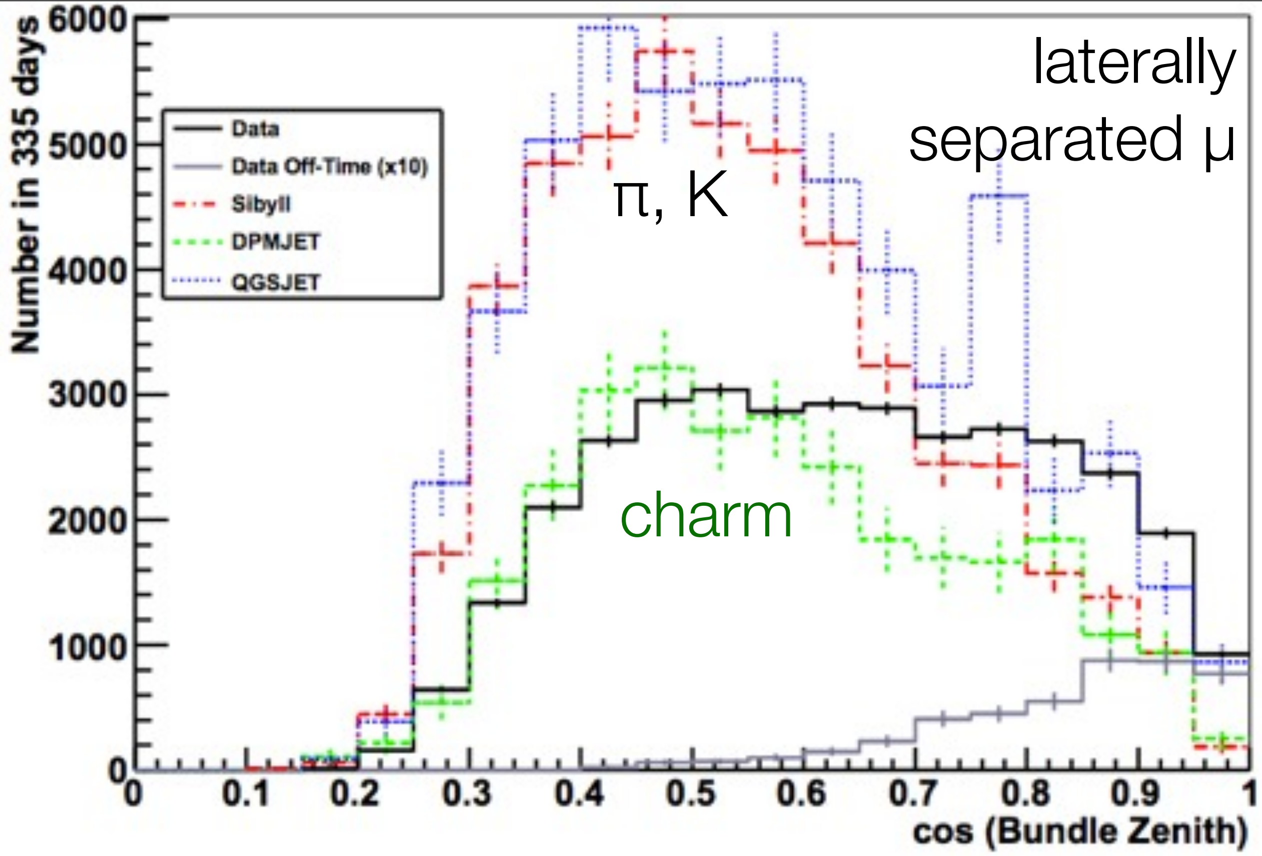
soft interaction (pT ≲ 2 GeV)
non perturbative

hard interaction (pT ≳ 2 GeV)
pQCD

IceCube Coll., PRD 87, 012005, 2013
arXiv:1208.2979

atmospheric neutrinos

charm and high p_T muons



IceCube Coll., PRD 87, 012005, 2013
arXiv:1208.2979

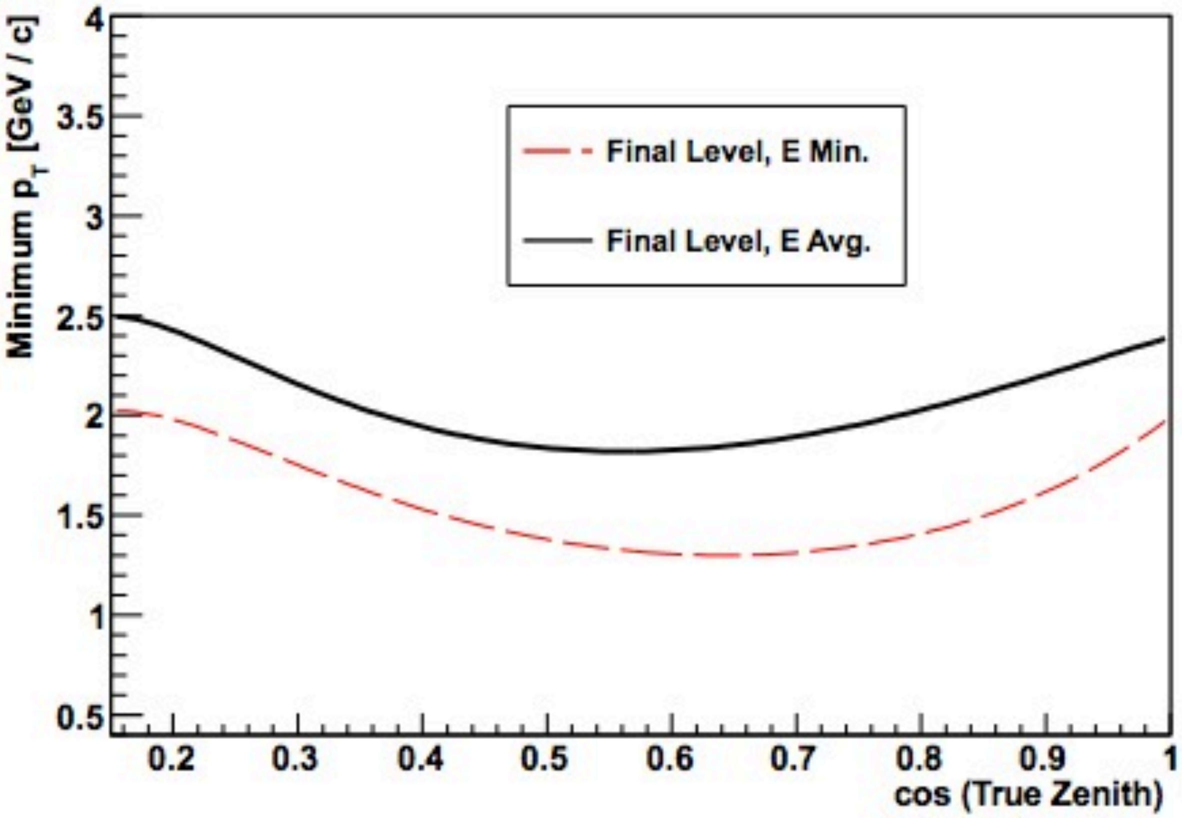
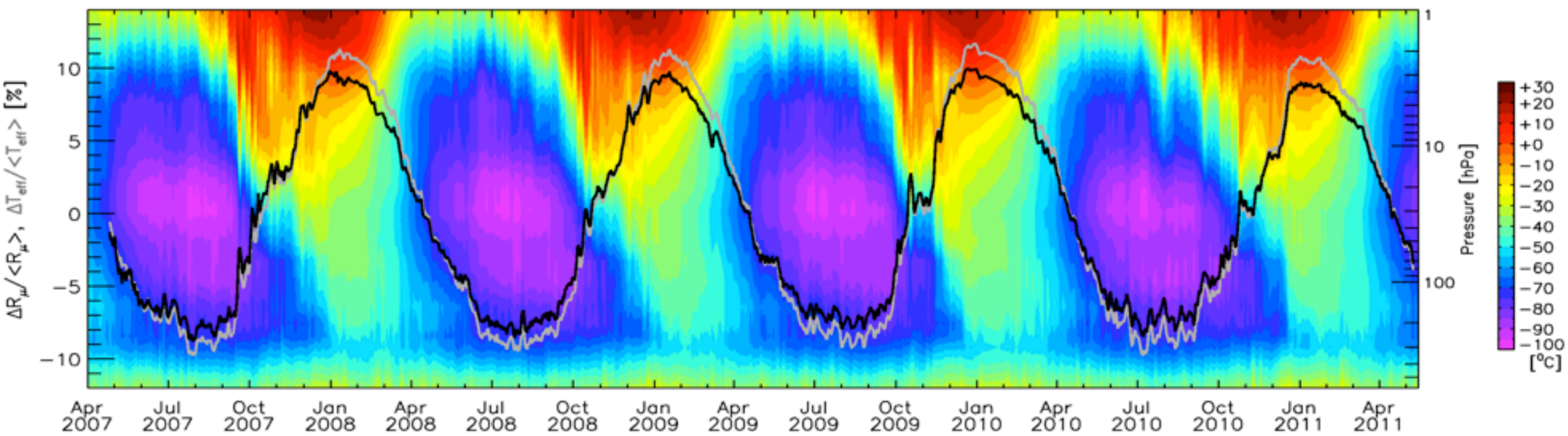
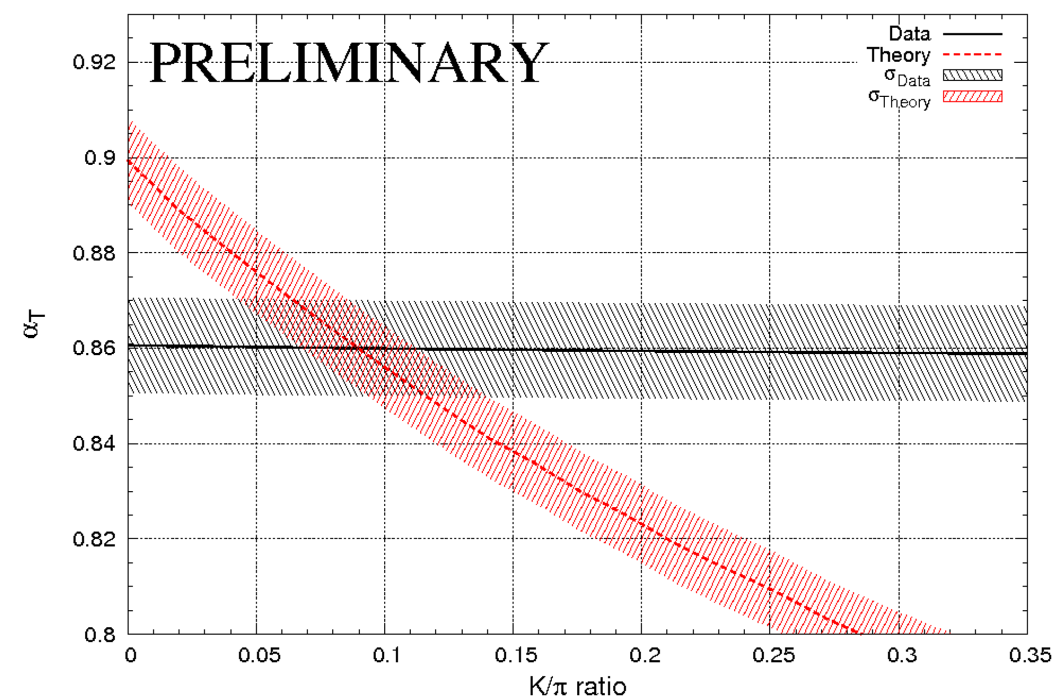


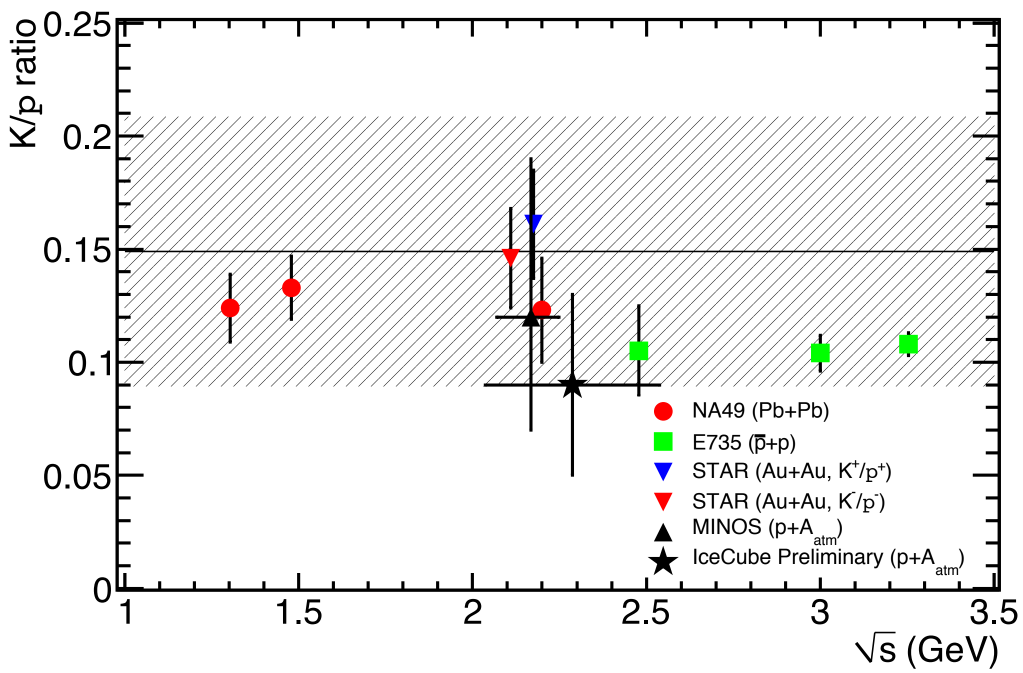
FIG. 14. (Color online). The minimum muon transverse momentum of DPMJET simulated shower events that pass all selection criteria for different energy parameterizations as a function of zenith angle. The interaction height comes from Fig. 1.

atmospheric neutrinos
π/K & μ seasonal variations

PD et al., ICRC 2011

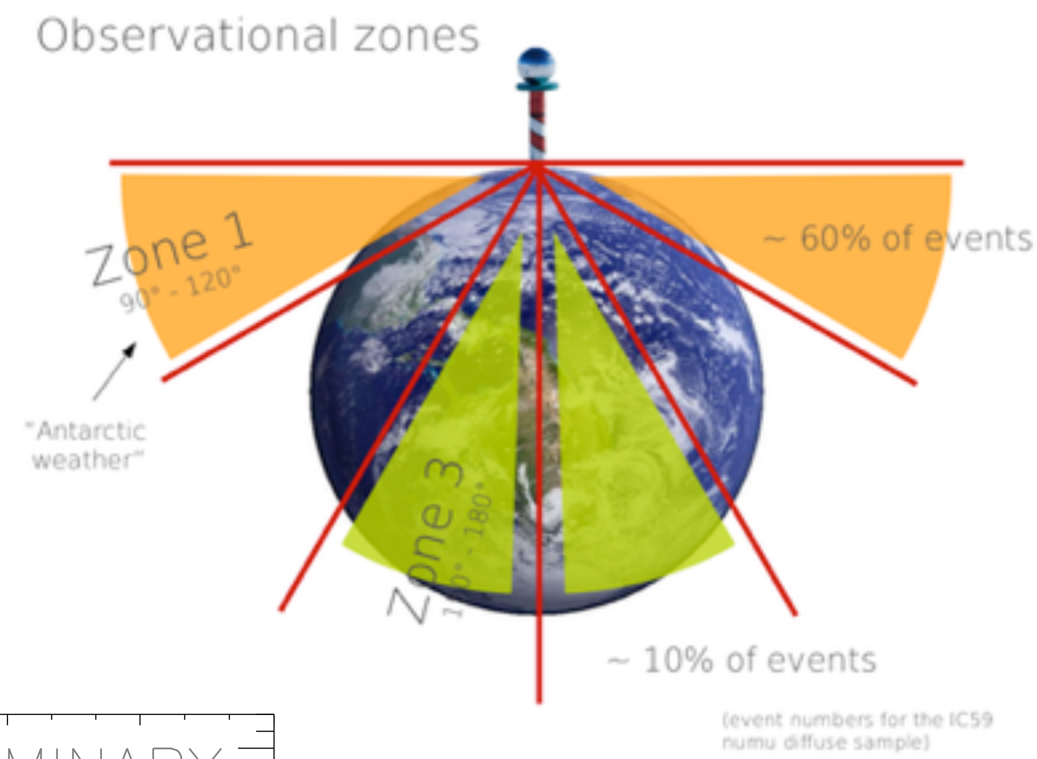






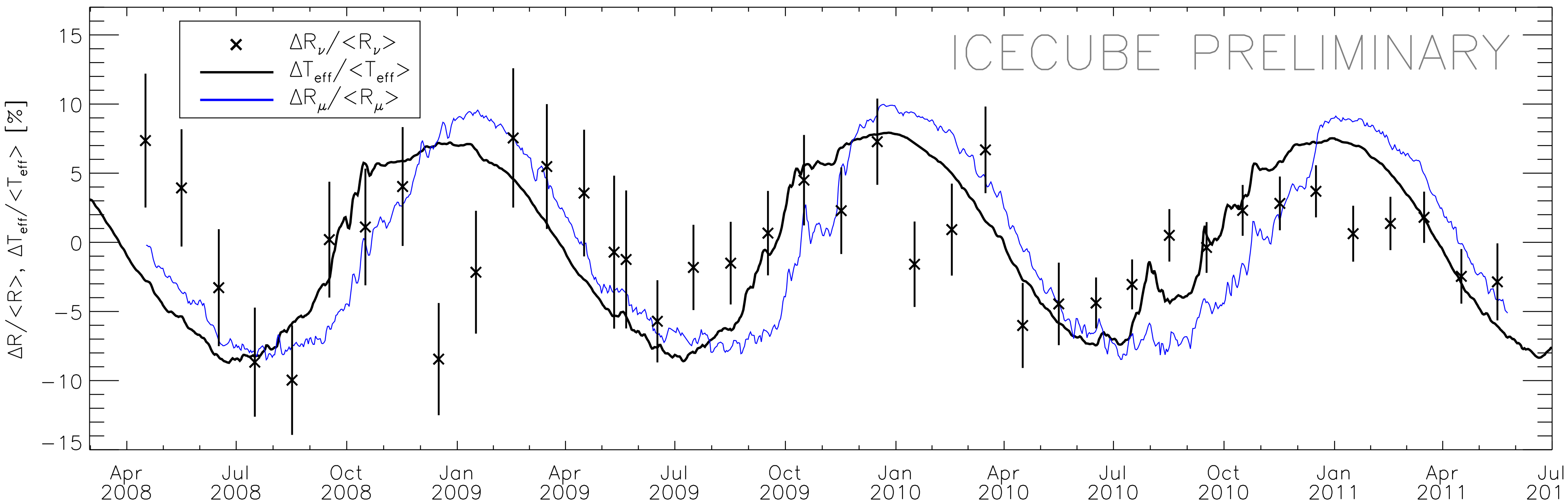
atmospheric neutrinos

charm & **ν** seasonal variations



effective temperature

$$T_{eff}(\theta) = \frac{\int dE_\nu \int dX \ A_{eff}(E_\nu,\theta)P(E_\nu,\theta,X)T(\theta,X)}{\int dE_\nu \int dX \ A_{eff}(E_\nu,\theta)P(E_\nu,\theta,X)}$$



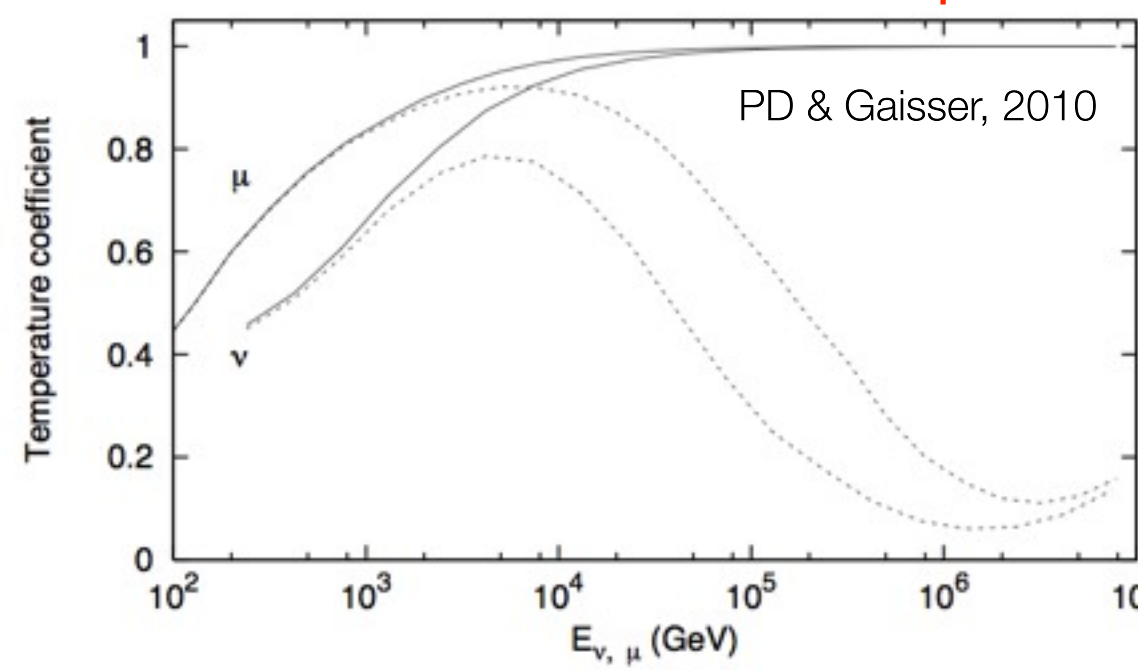
Tilav et al., ICRC 2009
PD et al., ICRC 2011
PD et al., ICRC 2013

seasonal variations decrease with prompt component

temperature coefficient

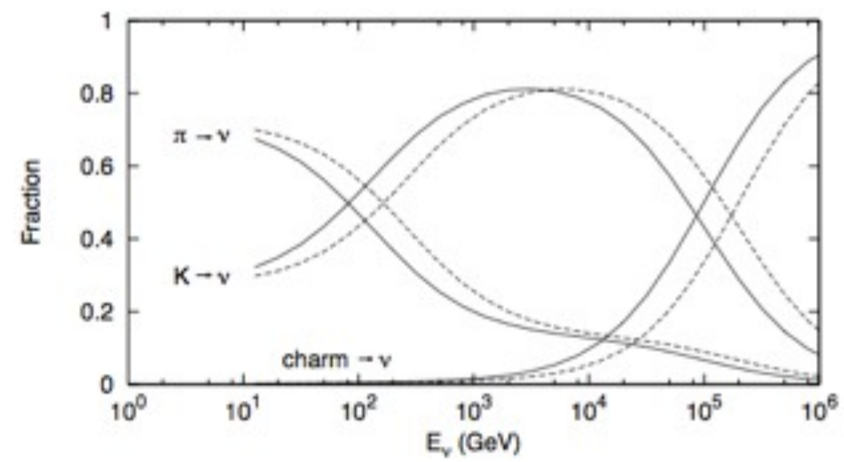
$$\alpha_T^{th}(\theta) = \frac{T \cdot \frac{\partial}{\partial T}\int dE_\nu \, \phi_\nu(E_\nu,\theta) A_{\rm eff}(E_\nu,\theta)}{\int dE_\nu \, \phi_\nu(E_\nu,\theta) A_{\rm eff}(E_\nu,\theta)}$$

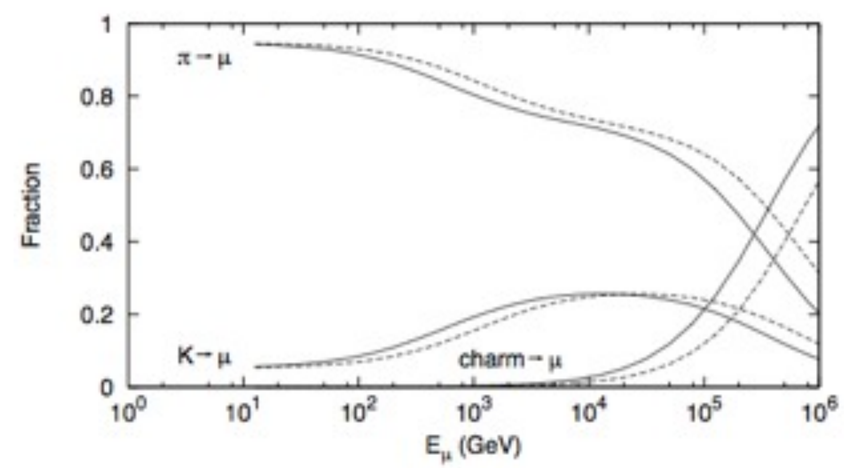
$$\frac{\Delta R_\nu}{\langle R_\nu \rangle} = \alpha_T^{exp} \frac{\Delta T_{\rm eff}}{\langle T_{\rm eff} \rangle}$$

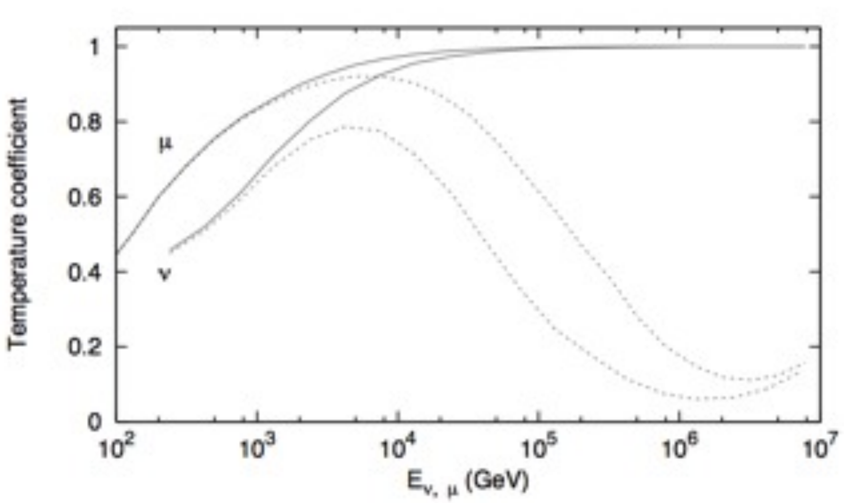


atmospheric neutrinos
π/K & ν seasonal variations

PD & Gaisser, 2010







$E_{\mu,\min}$	no charm		RQPM charm		ERS charm		int. charm	
	α	Rate	α	Rate	α	Rate	α	Rate
0.5	0.83	2050	0.82	2070	0.82	2050	0.82	2060
10	0.98	1.26	0.89	1.40	0.97	1.26	0.94	1.34
100	1.0	0.0025	0.53	0.0049	0.91	0.0028	0.71	0.0036

TABLE I: Correlation coefficients for muons with ($\theta \leq 30°$) for three levels of charm (energy in TeV; rate in Hz/km^2).

$E_{\nu,\min}$(TeV)	no charm		RQPM charm	
Zone 1	α	Events/yr	α	Events/yr
all	0.54	16000	0.52	17000
3	0.70	5900	0.62	6300
30	0.94	350	0.72	450
$E_{\nu,\min}$(TeV)	no charm		RQPM charm	
Zone 2	α	Events/yr	α	Events/yr
all	0.66	6000	0.62	6400
3	0.88	1230	0.75	1450
30	0.98	37	0.46	80
$E_{\nu,\min}$(TeV)	no charm		RQPM charm	
Zone 3	α	Events/yr	α	Events/yr
all	0.68	1650	0.64	1750
3	0.91	260	0.75	320
30	0.99	5.2	0.41	13

TABLE II: Correlation coefficients with and without charm for neutrinos in three zones of the atmosphere (see text).

PD et al., ICRC 2013

configuration	α_T^{exp}	χ^2/ndf	α_T^{th}
IC40	0.27±0.21	22.85/12	$0.557^{+0.008}_{-0.007}$
IC59	0.50±0.15	12.30/11	$0.518^{+0.008}_{-0.007}$
IC79	0.45±0.11	4.48/10	$0.489^{+0.007}_{-0.005}$